

2. ASTROPHYSICAL CONSTANTS

Table 2.1. Revised 1997 by D.E. Groom (LBNL) with the help of G.F. Smoot, M.S. Turner, and R.C. Willson. The figures in parentheses after some values give the one-standard deviation uncertainties in the last digit(s). While every effort has been made to obtain the most accurate current values of the listed quantities, the table does not represent a critical review or adjustment of the constants, and is not intended as a primary reference.

Quantity	Symbol, equation	Value	Reference
speed of light	c	299 792 458 m s ⁻¹	defined Ref. [1]
Newtonian gravitational constant	G_N	6.672 59(85) × 10 ⁻¹¹ m ³ kg ⁻¹ s ⁻²	Ref. [2]
astronomical unit	AU	1.495 978 706 6(2) × 10 ¹¹ m	Ref. [3,4]
tropical year (equinox to equinox) (1994)	yr	31 556 925.2 s	Ref. [3]
sidereal year (fixed star to fixed star) (1994)		31 558 149.8 s	Ref. [3]
mean sidereal day		23 ^h 56 ^m 04 ^s 090 53	Ref. [3]
Jansky	Jy	10 ⁻²⁶ W m ⁻² Hz ⁻¹	
Planck mass	$\sqrt{\hbar c/G_N}$	1.221 047(79) × 10 ¹⁹ GeV/c ² = 2.176 71(14) × 10 ⁻⁸ kg	uses Ref. [2]
parsec (1 AU/1 arc sec)	pc	3.085 677 580 7(4) × 10 ¹⁶ m = 3.262... ly	Ref. [5]
light year (deprecated unit)	ly	0.306 6... pc = 0.946 1... × 10 ¹⁶ m	
Schwarzschild radius of the Sun	$2G_N M_\odot/c^2$	2.953 250 08 km	Ref. [6]
solar mass	M_\odot	1.988 92(25) × 10 ³⁰ kg	Ref. [7]
solar luminosity	L_\odot	(3.846 ± 0.008) × 10 ²⁶ W	Ref. [8]
solar equatorial radius	R_\odot	6.96 × 10 ⁸ m	Ref. [3]
Earth equatorial radius	R_\oplus	6.378 140 × 10 ⁶ m	Ref. [3]
Earth mass	M_\oplus	5.973 70(76) × 10 ²⁴ kg	Ref. [9]
luminosity conversion	L	3.02 × 10 ²⁸ × 10 ^{-0.4 M_b} W (M _b = absolute bolometric magnitude = bolometric magnitude at 10 pc)	Ref. [10]
flux conversion	\mathcal{F}	2.52 × 10 ⁻⁸ × 10 ^{-0.4 m_b} W m ⁻² (m _b = apparent bolometric magnitude)	from above
v_\odot around center of Galaxy	Θ_\odot	220(20) km s ⁻¹	Ref. [11]
solar distance from galactic center	R_\odot	8.0(5) kpc	Ref. [12]
Hubble expansion rate [†]	H_0	100 h ₀ km s ⁻¹ Mpc ⁻¹ = h ₀ × (9.778 13 Gyr) ⁻¹	Ref. [13]
normalized Hubble expansion rate [†]	h_0	0.6 < h ₀ < 0.8	Ref. [14]
critical density of the universe [†]	$\rho_c = 3H_0^2/8\pi G_N$	2.775 366 27 × 10 ¹¹ h ₀ ² M _⊙ Mpc ⁻³ = 1.878 82(24) × 10 ⁻²⁹ h ₀ ² g cm ⁻³ = 1.053 94(13) × 10 ⁻⁵ h ₀ ² GeV cm ⁻³	
local disk density	ρ_{disk}	3–12 × 10 ⁻²⁴ g cm ⁻³ ≈ 2–7 GeV/c ² cm ⁻³	Ref. [15]
local halo density	ρ_{halo}	2–13 × 10 ⁻²⁵ g cm ⁻³ ≈ 0.1–0.7 GeV/c ² cm ⁻³	Ref. [16]
pressureless matter density of the universe [†]	$\Omega_M \equiv \rho_M/\rho_c$	0.2 < Ω _M < 1	Ref. [17]
scaled cosmological constant [†]	$\Omega_\Lambda = \Lambda c^2/3H_0^2$	-1 < Ω _Λ < 2	Ref. [18]
scale factor for cosmological constant [†]	$c^2/3H_0^2$	2.853 × 10 ⁵¹ h ₀ ⁻² m ²	
age of the universe [†]	t_0	11.5 + 1 ± 1.5 Gyr	Ref. [19]
	$\Omega_0 h_0^2$ for Λ = 0	≤ 2.4 for t ₀ ≥ 10 Gyr	Ref. [10]
		≤ 1 for t ₀ ≥ 10 Gyr, h ₀ > 0.4	Ref. [10]
		≤ 0.53 for t ₀ ≥ 10 Gyr, h ₀ > 0.6	Ref. [20]
cosmic background radiation (CBR) temperature [†]	T_0	2.728 ± 0.002 K	Ref. [21,22]
solar velocity with respect to CBR		369.3 ± 2.5 km s ⁻¹	Ref. [22,23]
energy density of CBR	ρ_γ	4.662 3 × 10 ⁻³⁴ (T/2.728) ⁴ g cm ⁻³ = 0.261 53 (T/2.728) ⁴ eV cm ⁻³	Ref. [10,22]
energy density of relativistic particles (CBR + ν)	ρ_R	7.838 8 × 10 ⁻³⁴ (T/2.728) ⁴ g cm ⁻³ = 0.439 72 (T/2.728) ⁴ eV cm ⁻³	Ref. [10,22]
number density of CBR photons	n_γ	411.87 (T/2.728) ³ cm ⁻³	Ref. [10,22]
entropy density/Boltzmann constant	s/k	2 899.3 (T/2.728) ³ cm ⁻³	Ref. [10]

[†] Subscript 0 indicates present-day values.

References:

1. B.W. Petley, *Nature* **303**, 373 (1983).
2. E.R. Cohen and B.N. Taylor, *Rev. Mod. Phys.* **59**, 1121 (1987).
The set of constants resulting from this adjustment has been recommended for international use by CODATA (Committee on Data for Science and Technology).
In the context of the scale dependence of field theoretic quantities, it should be remarked that absolute lab measurements of G_N have been performed only on scales of $10^{-1\pm 1}$ m.
3. *The Astronomical Almanac for the year 1994*, U.S. Government Printing Office, Washington, and Her Majesty's Stationary Office, London (1993). Where possible, the values as adjusted for the fitting of the ephemerides to all the observational data are used.
4. JPL Planetary Ephemerides, E. Myles Standish, Jr., private communication (1989).
5. 1 AU divided by $\pi/648\,000$; quoted error is from the JPL Planetary Ephemerides value of the AU [4].
6. Heliocentric gravitational constant from Ref. 3 times $2/c^2$. The given 9-place accuracy appears to be consistent with uncertainties in actually defining the earth's orbital parameters.
7. Obtained from the heliocentric gravitational constant [3] and G_N [2]. The error is the 128 ppm standard deviation of G_N .
8. 1996 mean total solar irradiance (TSI) = 1367.5 ± 2.7 [24]; the solar luminosity is $4\pi \times (1 \text{ AU})^2$ times this quantity. This value increased by 0.036% between the minima of solar cycles 21 and 22. It was modulated with an amplitude of 0.039% during solar cycle 21 [25].
Sackmann *et al.* [26] use TSI = $1370 \pm 2 \text{ W m}^{-2}$, but conclude that the solar luminosity ($L_\odot = 3.853 \times 10^{26} \text{ J s}^{-1}$) has an uncertainty of 1.5%. Their value is based on three 1977–83 papers, and they comment that the error is based on scatter among the reported values, which is substantially in excess of that expected from the individual quoted errors.
The conclusion of the 1971 review by Thekaekara and Drummond [27] ($1353 \pm 1\% \text{ W m}^{-2}$) is often quoted [28]. The conversion to luminosity is not given in the Thekaekara and Drummond paper, and we cannot exactly reproduce the solar luminosity given in Ref. 28.
Finally, a value based on the 1954 spectral curve due to Johnson [29] ($1395 \pm 1\% \text{ W m}^{-2}$, or $L_\odot = 3.92 \times 10^{26} \text{ J s}^{-1}$) has been used widely, and may be the basis for higher value of the solar luminosity and corresponding lower value of the solar absolute bolometric magnitude (4.72) still common in the literature [10].
9. Obtained from the geocentric gravitational constant [3] and G_N [2]. The error is the 128 ppm standard deviation of G_N .
10. E.W. Kolb and M.S. Turner, *The Early Universe*, Addison-Wesley (1990).
11. F.J. Kerr and D. Lynden-Bell, *Mon. Not. R. Astr. Soc.* **221**, 1023–1038 (1985). “On the basis of this review these [$R_\odot = 8.5 \pm 1.1 \text{ kpc}$ and $\Theta_\odot = 220 \pm 20 \text{ km s}^{-1}$] were adopted by resolution of IAU Commission 33 on 1985 November 21 at Delhi”.
12. M.J. Reid, *Annu. Rev. Astron. Astrophys.* **31**, 345–372 (1993). Note that Θ_\odot from the 1985 IAU Commission 33 recommendations is adopted in this review, although the new value for R_\odot is smaller.
13. Conversion using length of tropical year.
14. See the section on the Hubble Constant (Sec. 17 of this *Review*).
15. G. Gilmore, R.F.G. Wyse, and K. Kuijken, *Annu. Rev. Astron. Astrophys.* **27**, 555 (1989).
16. E.I. Gates, G. Gyuk, and M.S. Turner (*Astrophys. J.* **449**, L133 (1995)) find the local halo density to be $9.2_{-3.1}^{+3.8} \times 10^{-25} \text{ g cm}^{-3}$, but also comment that previously published estimates are in the range $1\text{--}10 \times 10^{-25} \text{ g cm}^{-3}$.
The value $0.3 \text{ GeV}/c^2$ has been taken as “standard” in several papers setting limits on WIMP mass limits, *e.g.* in M. Mori *et al.*, *Phys. Lett.* **B289**, 463 (1992).
17. As of April 1998 the consensus of observations seems to be $0.2 < \Omega_M < 0.5$, but systematic effects which raise the upper limit cannot be ruled out.
18. S.M. Carroll and W. H. Press, *Annu. Rev. Astron. Astrophys.* **30**, 499 (1992);
J. L. Tonry, in *Proc. Texas/PASCOS 92: Relativistic Astrophysics and Particle Cosmology*, ed. C.W. Akerlof and M. Srednicki (*Ann. NY Acad. Sci.* **688**, 113 (1993);
Work being reported as of April 1998 suggests a narrower range, possible excluding 0 in favor of a positive value.
19. B. Chaboyer, P. Demarque, P.J. Kernan, and L.M. Krauss, eprint **astro-ph/9706128 v3** (submitted to *Astrophys. J.*). The paper uses the recent Hipparcos parallax catalog to reanalyze globular cluster ages. The “+1” adds 1 Gy for the formation time of globular clusters.
20. J. Wells, private communication (1998).
21. D.J. Fixsen *et al.*, *Astrophys. J.* **473**, 576 (1996). Error quoted here is one standard deviation.
22. See the section on Cosmic Background Radiation (Sec. 19 of this *Review*).
23. C.H. Lineweaver *et al.*, *Astrophys. J.* **470**, 28 (1996). Dipole velocity is in the direction $(\ell, b) = (264^\circ.31 \pm 0^\circ.04 \pm 0^\circ.16, +48^\circ.05 \pm 0^\circ.02 \pm 0^\circ.09)$, or $(\alpha, \delta) = (11^{\text{h}}11^{\text{m}}57^{\text{s}} \pm -7^\circ.22 \pm 0^\circ.08)$ (JD2000).
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the 0.2% error estimate is from R.C. Willson, private correspondence (1998).
25. R.C. Willson and H.S. Hudson, *Nature* **332**, 810 (1988).
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K.R. Lang, *Astrophysical Data: Planets and Stars*, Springer-Verlag (1992).
29. F.S. Johnson, *J. Meteorol.* **11**, 431 (1954).
30. G.H. Jacoby *et al.*, *J. Astron. Soc. Pacific* **104**, 599–662 (1992).
31. J.P. Huchra, *Science* **256**, 321–325 (1992).