Citation: W.-M. Yao *et al.* (Particle Data Group), J. Phys. G **33**, 1 (2006) (URL: http://pdg.lbl.gov)

Neutrino Properties

INTRODUCTION TO THE NEUTRINO PROPERTIES LISTINGS

Revised May 2006 by P. Vogel (Caltech) and A. Piepke (University of Alabama).

The following Listings concern measurements of various properties of neutrinos. Nearly all of the measurements, all of which so far are upper limits, actually concern superpositions of the mass eigenstates ν_i, which are in turn related to the weak eigenstates ν_ℓ, via the neutrino mixing matrix

$$|\nu_\ell\rangle = \sum_i U_{\ell i}\,|\nu_i\rangle.$$

In the analogous case of quark mixing via the CKM matrix, the smallness of the off-diagonal terms (small mixing angles) permits a "dominant eigenstate" approximation. Previous editions of this *Review* had assumed that the dominant eigenstate paradigm applies to neutrinos as well. However, the present results of neutrino oscillation searches show that the mixing matrix contains two large mixing angles. We cannot, therefore, associate any particular state $|\nu_i\rangle$ with any particular lepton label e, μ or τ. Nevertheless, neutrinos are produced in weak decays with a definite lepton flavor, and are typically detected by the charged current weak interaction again associated with a specific lepton flavor. The Listings for the neutrino mass that follow are separated into the three associated charged-lepton categories. Other properties (mean lifetime, magnetic moment, charge, and charge radius) are no longer separated this way. If needed, the associated lepton flavor is reported in the footnotes.

Measured quantities (mass-squared, magnetic moments, mean lifetimes, *etc.*) all depend upon the mixing parameters $|U_{\ell i}|^2$, but to some extent also on experimental conditions (*e.g.*, on energy resolution). Most of these observables, in particular mass-squared, cannot distinguish between Dirac and Majorana neutrinos, and are unaffected by CP phases.

Direct neutrino mass measurements are usually based on the analysis of the kinematics of charged particles (leptons, pions) emitted together with neutrinos (flavor states) in various weak decays. The most sensitive neutrino mass measurement to date, involving electron type neutrinos, is based on fitting the shape of the beta spectrum. The quantity $\langle m_\beta^2 \rangle = \sum_i |U_{ei}|^2 m_{\nu_i}^2$ is determined or constrained, where the sum is over all mass eigenvalues m_{ν_i} that are too close together to be resolved experimentally. If the energy resolution is better than $\Delta m_{ij}^2 \equiv m_{\nu_i}^2 - m_{\nu_j}^2$, the corresponding heavier m_{ν_i} and mixing parameter could be determined by fitting the resulting spectral anomaly (step or kink).

A limit on $\langle m_\beta^2 \rangle$ implies an *upper* limit on the *minimum* value m_{min}^2 of $m_{\nu_i}^2$, independent of the mixing parameters U_{ei}: $m_{min}^2 \leq \langle m_\beta^2 \rangle$. However, if and when the study of neutrino oscillations provides us with the values of *all* neutrino mass-squared differences Δm_{ij}^2 and the mixing parameters $|U_{ei}|^2$, then the individual neutrino mass squares $m_{\nu_j}^2 = \langle m_\beta^2 \rangle - \sum_i |U_{ei}|^2 \Delta m_{ij}^2$ can be determined.

Leaving the yet unconfirmed LSND evidence aside, neutrino oscillation experiments using solar, reactor, atmospheric, and accelerator neutrinos can be described using two mass splittings and three mixing angles. Combined three neutrino analyses determine the squared mass differences and two of the mixing angles to within reasonable accuracy. For given $|\Delta m_{ij}^2|$, a limit on $\langle m_\beta^2 \rangle$ from beta decay defines an *upper* limit on the *maximum* value m_{max} of m_{ν_i}: $m_{max}^2 \leq \langle m_\beta^2 \rangle + \sum_{i<j} |\Delta m_{ij}^2|$. The analysis of the low energy beta decay of tritium, combined with the oscillation results, thus limits *all* neutrino masses. Traditionally experimental neutrino mass limits obtained from pion decay $\pi^+ \to \mu^+ + \nu_\mu$, or the shape of the spectrum of decay products of the τ lepton, did not distinguish between flavor and mass eigenstates. These results are reported as limits of the μ and τ based neutrino mass. After the determination of the $|\Delta m_{ij}^2|$'s, the corresponding neutrino mass limits are no longer competitive with those derived from low energy beta decays, with the proviso, however, that the oscillation searches, reported below, can be regarded as a reliable source of *all* $|\Delta m_{ij}^2|$ values.

The spread of arrival times of the neutrinos from SN1987A, coupled with the measured neutrino energies, provides a time-of-flight limit on a quantity similar to $\langle m_\beta \rangle \equiv \sqrt{\langle m_\beta^2 \rangle}$. This statement, clothed in various degrees of sophistication, has been the basis for a very large number of papers. The resulting limits, however, are no longer comparable with the limits from tritium beta decay.

Constraint on the sum of the neutrino masses can be obtained from the analysis of the cosmic microwave background anisotropy, combined with the galaxy redshift surveys and other data. These limits are reported in a separate table (Sum of Neutrino Masses, m_{tot}). Discussion concerning the model dependence of this limit is continuing.

$\overline{\nu}$ MASS (electron based)

Those limits given below are for the square root of $m_{\nu_e}^{2(\text{eff})} \equiv \sum_i |U_{ei}|^2 m_{\nu_i}^2$. Limits that come from the kinematics of ${}^3\text{H}\beta^-\overline{\nu}$ decay are the square roots of the limits for $m_{\nu_e}^{2(\text{eff})}$. Obtained from the measurements reported in the Listings for "$\overline{\nu}$ Mass Squared," below.

VALUE (eV)	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
< 2 OUR EVALUATION				
< 2.3	95	[1] KRAUS 05	SPEC	^{3}H β decay
< 2.5	95	[2] LOBASHEV 99	SPEC	^{3}H β decay
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<21.7	90	[3] ARNABOLDI 03A	BOLO	^{187}Re β-decay
< 5.7	95	[4] LOREDO 02	ASTR	SN1987A
< 2.8	95	[5] WEINHEIMER 99	SPEC	^{3}H β decay
< 4.35	95	[6] BELESEV 95	SPEC	^{3}H β decay
<12.4	95	[7] CHING 95	SPEC	^{3}Hβ decay
<92	95	[8] HIDDEMANN 95	SPEC	^{3}H β decay
$15 \; {+32 \atop -15}$		HIDDEMANN 95	SPEC	^{3}H β decay
<19.6	95	KERNAN 95	ASTR	SN 1987A
< 7.0	95	[9] STOEFFL 95	SPEC	^{3}H β decay
< 7.2	95	[10] WEINHEIMER 93	SPEC	^{3}H β decay
<11.7	95	[11] HOLZSCHUH 92B	SPEC	^{3}H β decay
<13.1	95	[12] KAWAKAMI 91	SPEC	^{3}H β decay
< 9.3	95	[13] ROBERTSON 91	SPEC	^{3}H β decay
<14	95	AVIGNONE 90	ASTR	SN 1987A
<16		SPERGEL 88	ASTR	SN 1987A
17 to 40		[14] BORIS 87	SPEC	^{3}Hβ decay

[1] KRAUS 05 is a continuation of the work reported in WEINHEIMER 99. This result represents the final analysis of data taken from 1997 to 2001. Various sources of systematic uncertainties have been identified and quantified. The background has been reduced compared to the initial running period. A spectral anomaly at the endpoint, reported in LOBASHEV 99, was not observed.

[2] LOBASHEV 99 report a new measurement which continues the work reported in BELESEV 95. This limit depends on phenomenological fit parameters used to derive their best fit to m_ν^2, making unambiguous interpretation difficult. See the footnote under "$\overline{\nu}$ Mass Squared."

[3] ARNABOLDI 03A *etal.* report kinematical neutrino mass limit using β-decay of ^{187}Re. Bolometric $AgReO_4$ micro-calorimeters are used. Mass bound is substantially weaker than those derived from tritium β-decays but has different systematic uncertainties.

[4] LOREDO 02 updates LOREDO 89.

[5] WEINHEIMER 99 presents two analyses which exclude the spectral anomaly and result in an acceptable m_ν^2. We report the most conservative limit, but the other is nearly the same. See the footnote under "$\overline{\nu}$ Mass Squared."

[6] BELESEV 95 (Moscow) use an integral electrostatic spectrometer with adiabatic magnetic collimation and a gaseous tritium sources. A fit to a normal Kurie plot above 18300–18350 eV (to avoid a low-energy anomaly) plus a monochromatic line 7–15 eV below the endpoint yields $m_\nu^2 = -4.1 \pm 10.9$ eV^2, leading to this Bayesian limit.

[7] CHING 95 quotes results previously given by SUN 93; no experimental details are given. A possible explanation for consistently negative values of m_ν^2 is given.

[8] HIDDEMANN 95 (Munich) experiment uses atomic tritium embedded in a metal-dioxide lattice. Bayesian limit calculated from the weighted mean $m_\nu^2 = 221 \pm 4244$ eV^2 from the two runs listed below.

[9] STOEFFL 95 (LLNL) result is the Bayesian limit obtained from the m_ν^2 errors given below but with m_ν^2 set equal to 0. The anomalous endpoint accumulation leads to a value of m_ν^2 which is negative by more than 5 standard deviations.

[10] WEINHEIMER 93 (Mainz) is a measurement of the endpoint of the tritium β spectrum using an electrostatic spectrometer with a magnetic guiding field. The source is molecular tritium frozen onto an aluminum substrate.

[11] HOLZSCHUH 92B (Zurich) result is obtained from the measurement $m_\nu^2 = -24 \pm 48 \pm 61$ (1σ errors), in eV^2, using the PDG prescription for conversion to a limit in m_ν.

[12] KAWAKAMI 91 (Tokyo) experiment uses tritium-labeled arachidic acid. This result is the Bayesian limit obtained from the m_ν^2 limit with the errors combined in quadrature. This was also done in ROBERTSON 91, although the authors report a different procedure.

[13] ROBERTSON 91 (LANL) experiment uses gaseous molecular tritium. The result is in strong disagreement with the earlier claims by the ITEP group [LUBIMOV 80, BORIS 87 (+ BORIS 88 erratum)] that m_ν lies between 17 and 40 eV. However, the probability of a positive m^2 is only 3% if statistical and systematic error are combined in quadrature.

[14] See also comment in BORIS 87B and erratum in BORIS 88.

$\overline{\nu}$ MASS SQUARED (electron based)

Given troubling systematics which result in improbably negative estimators of $m_{\nu_e}^{2(\text{eff})} \equiv \sum_i |U_{ei}|^2\ m_{\nu_i}^2$, in many experiments, we use only KRAUS 05 and LOBASHEV 99 for our average.

VALUE (eV^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
− 1.1± 2.4 OUR AVERAGE					
− 0.6± 2.2± 2.1		[15] KRAUS	05	SPEC	^{3}H β decay
− 1.9± 3.4± 2.2		[16] LOBASHEV	99	SPEC	^{3}H β decay
• • • We do not use the following data for averages, fits, limits, etc. • • •					
− 3.7± 5.3± 2.1		[17] WEINHEIMER	99	SPEC	^{3}H β decay
− 22 ± 4.8		[18] BELESEV	95	SPEC	^{3}H β decay
129 ±6010		[19] HIDDEMANN	95	SPEC	^{3}H β decay
313 ±5994		[19] HIDDEMANN	95	SPEC	^{3}H β decay
−130 ± 20 ±15	95	[20] STOEFFL	95	SPEC	^{3}H β decay
− 31 ± 75 ±48		[21] SUN	93	SPEC	^{3}Hβ decay
− 39 ± 34 ±15		[22] WEINHEIMER	93	SPEC	^{3}H β decay
− 24 ± 48 ±61		[23] HOLZSCHUH	92B	SPEC	^{3}H β decay
− 65 ± 85 ±65		[24] KAWAKAMI	91	SPEC	^{3}H β decay
−147 ± 68 ±41		[25] ROBERTSON	91	SPEC	^{3}H β decay

[15] KRAUS 05 is a continuation of the work reported in WEINHEIMER 99. This result represents the final analysis of data taken from 1997 to 2001. Problems with significantly negative squared neutrino masses, observed in some earlier experiments, have been resolved in this work.

[16] LOBASHEV 99 report a new measurement which continues the work reported in BELESEV 95. The data were corrected for electron trapping effects in the source, eliminating the dependence of the fitted neutrino mass on the fit interval. The analysis assuming a pure beta spectrum yields significantly negative fitted $m_\nu^2 \approx -(20\text{–}10)$ eV^2. This problem is attributed to a discrete spectral anomaly of about 6×10^{-11} intensity with a time-dependent energy of 5–15 eV below the endpoint. The data analysis accounts for this anomaly by introducing two extra phenomenological fit parameters resulting in a best fit of $m_\nu^2 = -1.9 \pm 3.4 \pm 2.2\,\text{eV}^2$ which is used to derive a neutrino mass limit. However, the introduction of phenomenological fit parameters which are correlated with the derived m_ν^2 limit makes unambiguous interpretation of this result difficult.

[17] WEINHEIMER 99 is a continuation of the work reported in WEINHEIMER 93 . Using a lower temperature of the frozen tritium source eliminated the dewetting of the T_2 film, which introduced a dependence of the fitted neutrino mass on the fit interval in the earlier work. An indication for a spectral anomaly reported in LOBASHEV 99 has been seen, but its time dependence does not agree with LOBASHEV 99. Two analyses, which exclude the spectral anomaly either by choice of the analysis interval or by using a particular data set which does not exhibit the anomaly, result in acceptable m_ν^2 fits and are used to derive the neutrino mass limit published by the authors. We list the most conservative of the two.

[18] BELESEV 95 (Moscow) use an integral electrostatic spectrometer with adiabatic magnetic collimation and a gaseous tritium sources. This value comes from a fit to a normal Kurie plot above 18300–18350 eV (to avoid a low-energy anomaly), including the effects of an apparent peak 7–15 eV below the endpoint.

[19] HIDDEMANN 95 (Munich) experiment uses atomic tritium embedded in a metal-dioxide lattice. They quote measurements from two data sets.

[20] STOEFFL 95 (LLNL) uses a gaseous source of molecular tritium. An anomalous pileup of events at the endpoint leads to the negative value for m_ν^2. The authors acknowledge that "the negative value for the best fit of m_ν^2 has no physical meaning" and discuss possible explanations for this effect.

[21] SUN 93 uses a tritiated hydrocarbon source. See also CHING 95.

[22] WEINHEIMER 93 (Mainz) is a measurement of the endpoint of the tritium β spectrum using an electrostatic spectrometer with a magnetic guiding field. The source is molecular tritium frozen onto an aluminum substrate.

[23] HOLZSCHUH 92B (Zurich) source is a monolayer of tritiated hydrocarbon.

[24] KAWAKAMI 91 (Tokyo) experiment uses tritium-labeled arachidic acid.

[25] ROBERTSON 91 (LANL) experiment uses gaseous molecular tritium. The result is in strong disagreement with the earlier claims by the ITEP group [LUBIMOV 80, BORIS 87 (+ BORIS 88 erratum)] that m_ν lies between 17 and 40 eV. However, the probability of a positive m_ν^2 is only 3% if statistical and systematic error are combined in quadrature.

ν MASS (electron based)

These are measurement of m_ν (in contrast to $m_{\overline{\nu}}$, given above). The masses can be different for a Dirac neutrino in the absence of *CPT* invariance. The possible distinction between ν and $\overline{\nu}$ properties is usually ignored elsewhere in these Listings.

VALUE (eV)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
<460	68	YASUMI	94	CNTR	^{163}Ho decay
<225	95	SPRINGER	87	CNTR	^{163}Ho decay

ν MASS (muon based)

Limits given below are for the square root of $m^{2(\text{eff})}_{\nu_\mu} \equiv \sum_i |U_{\mu i}|^2\, m^2_{\nu_i}$.

In some of the COSM papers listed below, the authors did not distinguish between weak and mass eigenstates.

OUR EVALUATION is based on OUR AVERAGE for the $\pi^\pm$ mass and the ASSAMAGAN 96 value for the muon momentum for the π^+ decay at rest. The limit is calculated using the unified classical analysis of FELDMAN 98 for a Gaussian distribution near a physical boundary. WARNING: since $m^{2(\text{eff})}_{\nu_\mu}$ is calculated from the differences of large numbers, it and the corresponding limits are extraordinarily sensitive to small changes in the pion mass, the decay muon momentum, and their errors. For example, the limits obtained using JECKELMANN 94, LENZ 98, and the weighted averages are 0.15, 0.29, and 0.19 MeV, respectively.

VALUE (MeV)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
<0.19 (CL = 90%) OUR EVALUATION					
<0.17	90	[26] ASSAMAGAN	96	SPEC	$m^2_\nu = -0.016 \pm 0.023$
• • • We do not use the following data for averages, fits, limits, etc. • • •					
<0.15		[27] DOLGOV	95	COSM	Nucleosynthesis
<0.48		[28] ENQVIST	93	COSM	Nucleosynthesis
<0.3		[29] FULLER	91	COSM	Nucleosynthesis
<0.42		[29] LAM	91	COSM	Nucleosynthesis
<0.50	90	[30] ANDERHUB	82	SPEC	$m^2_\nu = -0.14 \pm 0.20$
<0.65	90	CLARK	74	ASPK	$K_{\mu 3}$ decay

[26] ASSAMAGAN 96 measurement of p_μ from $\pi^+ \to \mu^+ \nu$ at rest combined with JECKELMANN 94 Solution B pion mass yields $m^2_\nu = -0.016 \pm 0.023$ with corresponding Bayesian limit listed above. If Solution A is used, $m^2_\nu = -0.143 \pm 0.024$ MeV^2. Replaces ASSAMAGAN 94.

[27] DOLGOV 95 removes earlier assumptions (DOLGOV 93) about thermal equilibrium below T_{QCD} for wrong-helicity Dirac neutrinos (ENQVIST 93, FULLER 91) to set more stringent limits.

[28] ENQVIST 93 bases limit on the fact that thermalized wrong-helicity Dirac neutrinos would speed up expansion of early universe, thus reducing the primordial abundance. FULLER 91 exploits the same mechanism but in the older calculation obtains a larger production rate for these states, and hence a lower limit. Neutrino lifetime assumed to exceed nucleosynthesis time, ~ 1 s.

[29] Assumes neutrino lifetime >1 s. For Dirac neutrinos only. See also ENQVIST 93.

[30] ANDERHUB 82 kinematics is insensitive to the pion mass.

ν MASS (tau based)

The limits given below are the square roots of limits for $m^{2(\text{eff})}_{\nu_\tau} \equiv \sum_i |\text{U}_{\tau i}|^2\ m^2_{\nu_i}$.

In some of the ASTR and COSM papers listed below, the authors did not distinguish between weak and mass eigenstates.

VALUE (MeV)	*CL%*	*EVTS*	*DOCUMENT ID*	*TECN*	*COMMENT*
< 18.2	95		[31] BARATE 98F	ALEP	1991–1995 LEP runs
• • • We do not use the following data for averages, fits, limits, etc. • • •					
$<$ 28	95		[32] ATHANAS 00	CLEO	E^{ee}_{cm}= 10.6 GeV
$<$ 27.6	95		[33] ACKERSTAFF 98T	OPAL	1990–1995 LEP runs
$<$ 30	95	473	[34] AMMAR 98	CLEO	E^{ee}_{cm} = 10.6 GeV
$<$ 60	95		[35] ANASTASSOV 97	CLEO	E^{ee}_{cm}= 10.6 GeV
$<$ 0.37 or $>$22			[36] FIELDS 97	COSM	Nucleosynthesis
$<$ 68	95		[37] SWAIN 97	THEO	m_τ, τ_τ, τ partial widths
$<$ 29.9	95		[38] ALEXANDER 96M	OPAL	1990–1994 LEP runs
$<$149			[39] BOTTINO 96	THEO	π, μ, τ leptonic decays
$<$1 or $>$25			[40] HANNESTAD 96C	COSM	Nucleosynthesis
$<$ 71	95		[41] SOBIE 96	THEO	m_τ, τ_τ, B($\tau^- \to e^- \overline{\nu}_e \nu_\tau$)
$<$ 24	95	25	[42] BUSKULIC 95H	ALEP	1991–1993 LEP runs
$<$ 0.19			[43] DOLGOV 95	COSM	Nucleosynthesis
$<$ 3			[44] SIGL 95	ASTR	SN 1987A
$<$ 0.4 or $>$ 30			[45] DODELSON 94	COSM	Nucleosynthesis
$<$ 0.1 or $>$ 50			[46] KAWASAKI 94	COSM	Nucleosynthesis
155–225			[47] PERES 94	THEO	π,K,μ,τ weak decays
$<$ 32.6	95	113	[48] CINABRO 93	CLEO	$E^{ee}_{\text{cm}} \approx$ 10.6 GeV
$<$ 0.3 or $>$ 35			[49] DOLGOV 93	COSM	Nucleosynthesis
$<$ 0.74			[50] ENQVIST 93	COSM	Nucleosynthesis
$<$ 31	95	19	[51] ALBRECHT 92M	ARG	E^{ee}_{cm}= 9.4–10.6 GeV
$<$ 0.3			[52] FULLER 91	COSM	Nucleosynthesis
$<$ 0.5 or $>$ 25			[53] KOLB 91	COSM	Nucleosynthesis
$<$ 0.42			[52] LAM 91	COSM	Nucleosynthesis

[31] BARATE 98F result based on kinematics of 2939 $\tau^- \to 2\pi^-\pi^+\nu_\tau$ and 52 $\tau^- \to 3\pi^-2\pi^+(\pi^0)\nu_\tau$ decays. If possible 2.5% excited a_1 decay is included in 3-prong sample analysis, limit increases to 19.2 MeV.

[32] ATHANAS 00 bound comes from analysis of $\tau^- \to \pi^-\pi^+\pi^-\pi^0\nu_\tau$ decays.

[33] ACKERSTAFF 98T use $\tau \to 5\pi^\pm\nu_\tau$ decays to obtain a limit of 43.2 MeV (95%CL). They combine this with ALEXANDER 96M value using $\tau \to 3h^\pm\nu_\tau$ decays to obtain quoted limit.

[34] AMMAR 98 limit comes from analysis of $\tau^- \to 3\pi^-2\pi^+\nu_\tau$ and $\tau^- \to 2\pi^-\pi^+2\pi^0\nu_\tau$ decay modes.

[35] ANASTASSOV 97 derive limit by comparing their m_τ measurement (which depends on m_{ν_τ}) to BAI 96 m_τ threshold measurement.

[36] FIELDS 97 limit for a Dirac neutrino. For a Majorana neutrino the mass region < 0.93 or >31 MeV is excluded. These bounds assume $N_\nu <4$ from nucleosynthesis; a wider excluded region occurs with a smaller N_ν upper limit.

[37] SWAIN 97 derive their limit from the Standard Model relationships between the tau mass, lifetime, branching fractions for $\tau^- \rightarrow e^- \overline{\nu}_e \nu_\tau$, $\tau^- \rightarrow \mu^- \overline{\nu}_\mu \nu_\tau$, $\tau^- \rightarrow \pi^- \nu_\tau$, and $\tau^- \rightarrow K^- \nu_\tau$, and the muon mass and lifetime by assuming lepton universality and using world average values. Limit is reduced to 48 MeV when the CLEO τ mass measurement (BALEST 93) is included; see CLEO's more recent m_{ν_τ} limit (ANASTASSOV 97). Consideration of mixing with a fourth generation heavy neutrino yields $\sin^2\theta_L < 0.016$ (95%CL).

[38] ALEXANDER 96M bound comes from analyses of $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$ and $\tau^- \rightarrow h^- h^- h^+ \nu_\tau$ decays.

[39] BOTTINO 96 assumes three generations of neutrinos with mixing, finds consistency with massless neutrinos with no mixing based on 1995 data for masses, lifetimes, and leptonic partial widths.

[40] HANNESTAD 96C limit is on the mass of a Majorana neutrino. This bound assumes $N_\nu < 4$ from nucleosynthesis. A wider excluded region occurs with a smaller N_ν upper limit. This paper is the corrected version of HANNESTAD 96; see the erratum: HANNESTAD 96B.

[41] SOBIE 96 derive their limit from the Standard Model relationship between the tau mass, lifetime, and leptonic branching fraction, and the muon mass and lifetime, by assuming lepton universality and using world average values.

[42] BUSKULIC 95H bound comes from a two-dimensional fit of the visible energy and invariant mass distribution of $\tau \rightarrow 5\pi(\pi^0)\nu_\tau$ decays. Replaced by BARATE 98F.

[43] DOLGOV 95 removes earlier assumptions (DOLGOV 93) about thermal equilibrium below T_{QCD} for wrong-helicity Dirac neutrinos (ENQVIST 93, FULLER 91) to set more stringent limits. DOLGOV 96 argues that a possible window near 20 MeV is excluded.

[44] SIGL 95 exclude massive Dirac or Majorana neutrinos with lifetimes between 10^{-3} and 10^8 seconds if the decay products are predominantly γ or e^+e^-.

[45] DODELSON 94 calculate constraints on ν_τ mass and lifetime from nucleosynthesis for 4 generic decay modes. Limits depend strongly on decay mode. Quoted limit is valid for all decay modes of Majorana neutrinos with lifetime greater than about 300 s. For Dirac neutrinos limits change to < 0.3 or > 33.

[46] KAWASAKI 94 excluded region is for Majorana neutrino with lifetime $>$1000 s. Other limits are given as a function of ν_τ lifetime for decays of the type $\nu_\tau \rightarrow \nu_\mu \phi$ where ϕ is a Nambu-Goldstone boson.

[47] PERES 94 used PDG 92 values for parameters to obtain a value consistent with mixing. Reexamination by BOTTINO 96 which included radiative corrections and 1995 PDG parameters resulted in two allowed regions, $m_3 < 70$ MeV and 140 MeV $m_3 < 149$ MeV.

[48] CINABRO 93 bound comes from analysis of $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$ and $\tau^- \rightarrow 2\pi^- \pi^+ 2\pi^0 \nu_\tau$ decay modes.

[49] DOLGOV 93 assumes neutrino lifetime $>$100 s. For Majorana neutrinos, the low mass limit is 0.5 MeV. KAWANO 92 points out that these bounds can be overcome for a Dirac neutrino if it possesses a magnetic moment. See also DOLGOV 96.

[50] ENQVIST 93 bases limit on the fact that thermalized wrong-helicity Dirac neutrinos would speed up expansion of early universe, thus reducing the primordial abundance. FULLER 91 exploits the same mechanism but in the older calculation obtains a larger

production rate for these states, and hence a lower limit. Neutrino lifetime assumed to exceed nucleosynthesis time, ~ 1 s.

[51] ALBRECHT 92M reports measurement of a slightly lower τ mass, which has the effect of reducing the ν_τ mass reported in ALBRECHT 88B. Bound is from analysis of $\tau^- \rightarrow 3\pi^- 2\pi^+ \nu_\tau$ mode.

[52] Assumes neutrino lifetime >1 s. For Dirac neutrinos. See also ENQVIST 93.

[53] KOLB 91 exclusion region is for Dirac neutrino with lifetime >1 s; other limits are given.

Revised April 1998 by K.A. Olive (University of Minnesota).

The limits on low mass ($m_\nu \lesssim 1$ MeV) neutrinos apply to $m_{\rm tot}$ given by

$$m_{\rm tot} = \sum_\nu (g_\nu/2) m_\nu \ ,$$

where g_ν is the number of spin degrees of freedom for ν plus $\overline{\nu}$: $g_\nu = 4$ for neutrinos with Dirac masses; $g_\nu = 2$ for Majorana neutrinos. Stable neutrinos in this mass range make a contribution to the total energy density of the Universe which is given by

$$\rho_\nu = m_{\rm tot} n_\nu = m_{\rm tot}(3/11) n_\gamma \ ,$$

where the factor 3/11 is the ratio of (light) neutrinos to photons. Writing $\Omega_\nu = \rho_\nu/\rho_c$, where ρ_c is the critical energy density of the Universe, and using $n_\gamma = 412\ {\rm cm}^{-3}$, we have

$$\Omega_\nu h^2 = m_{\rm tot}/(94\ {\rm eV}) \ .$$

Therefore, a limit on $\Omega_\nu h^2$ such as $\Omega_\nu h^2 < 0.25$ gives the limit

$$m_{\rm tot} < 24\ {\rm eV} \ .$$

The limits on high mass ($m_\nu > 1$ MeV) neutrinos apply separately to each neutrino type.

SUM OF THE NEUTRINO MASSES, m_{tot}

(Defined in the above note), of effectively stable neutrinos (i.e., those with mean lives greater than or equal to the age of the universe). These papers assumed Dirac neutrinos. When necessary, we have generalized the results reported so they apply to m_{tot}. For other limits, see SZALAY 76, VYSOTSKY 77, BERNSTEIN 81, FREESE 84, SCHRAMM 84, and COWSIK 85.

VALUE (eV)	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$<$ 2.0	[54] ICHIKAWA	05	COSM	
$<$ 0.75	[55] BARGER	04	COSM	
$<$ 1.0	[56] CROTTY	04	COSM	
$<$ 1.0	[57] HANNESTAD	03B	COSM	
$<$ 0.7	[58] SPERGEL	03	COSM	WMAP
$<$ 1.8	[59] ELGAROY	02	ASTR	2dF Galaxy Redshift Survey
$<$ 0.9	[60] LEWIS	02	COSM	
$<$ 4.2	[61] WANG	02	COSM	CMB
$<$ 2.7	[62] FUKUGITA	00	COSM	
$<$ 5.5	[63] CROFT	99	ASTR	Ly α power spec
$<$180	SZALAY	74	COSM	
$<$132	COWSIK	72	COSM	
$<$280	MARX	72	COSM	
$<$400	GERSHTEIN	66	COSM	

[54] Constrains the total mass of neutrinos from the CMB experiments alone, assuming ΛCDM Universe.

[55] Constrains the total mass of neutrinos from the power spectrum of fluctuations derived from the Sloan Digital Sky Survey and the 2dF galaxy redshift survey, WMAP and 27 other CMB experiments and measurements by the HST Key project.

[56] Constrains the total mass of neutrinos from the power spectrum of fluctuations derived from the Sloan Digital Sky Survey, the 2dF galaxy redshift survey, WMAP and ACBAR. The limit is strengthened to 0.6 eV when measurements by the HST Key project and supernovae data are included.

[57] Constrains the fractional contribution of neutrinos to the total matter density in the Universe from WMAP data combined with other CMB measurements, the 2dfGRS data, HST data, and SN1a data.

[58] Constrains the fractional contribution of neutrinos to the total matter density in the Universe from WMAP data combined with other CMB measurements, the 2dfGRS data, and Lyman α data. The limit does not noticeably change if the Lyman α data are not used.

[59] ELGAROY 02 constrains the fractional contribution of neutrinos to the total matter density in the Universe from the power spectrum of fluctuations derived from the 2 Degree Field Galaxy Redshift Survey. Assumes $\Omega_{\text{matter}} < 0.5$ and a spectral index of 1.0. Limit softens to $m_\nu < 2.2$ eV for n=1.0 ± 0.1.

[60] LEWIS 02 constrains the total mass of neutrinos from the power spectrum of fluctuations derived from the CMB, HST Key project, 2dF galaxy redshift survey, supernovae type Ia, and BBN.

[61] WANG 02 constrains the total mass of neutrinos from the power spectrum of fluctuations derived from the CMB and other cosmological data sets such as galaxy clustering and the Lyman α forest.

[62] FUKUGITA 00 is a limit on neutrino masses from structure formation. The constraint is based on the clustering scale σ_8 and the COBE normalization and leads to a conservative limit of 0.9 eV assuming 3 nearly degenerate neutrinos. The quoted limit is on the sum of the light neutrino masses.

[63] CROFT 99 result based on the power spectrum of the Ly α forest. If $\Omega_{\text{matter}} < 0.5$, the limit is improved to $m_\nu < 2.4$ (Ω_{matter}/0.17–1) eV.

Limits on MASSES of Light Stable Right-Handed ν (with necessarily suppressed interaction strengths)

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<100–200	[64] OLIVE	82	COSM	Dirac ν
<200–2000	[64] OLIVE	82	COSM	Majorana ν

[64] Depending on interaction strength G_R where $G_R < G_F$.

Limits on MASSES of Heavy Stable Right-Handed ν (with necessarily suppressed interaction strengths)

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
> 10	[65] OLIVE	82	COSM	G_R/G_F <0.1
>100	[65] OLIVE	82	COSM	G_R/G_F <0.01

[65] These results apply to heavy Majorana neutrinos and are summarized by the equation: m_ν >1.2 GeV (G_F/G_R). The bound saturates, and if G_R is too small no mass range is allowed.

ν CHARGE

VALUE (units: electron charge)	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$<2 \times 10^{-14}$	[66] RAFFELT	99	ASTR	Red giant luminosity
$<6 \times 10^{-14}$	[67] RAFFELT	99	ASTR	Solar cooling
$<4 \times 10^{-4}$	[68] BABU	94	RVUE	BEBC beam dump
$<3 \times 10^{-4}$	[69] DAVIDSON	91	RVUE	SLAC electron beam dump
$<2 \times 10^{-15}$	[70] BARBIELLINI	87	ASTR	SN 1987A
$<1 \times 10^{-13}$	[71] BERNSTEIN	63	ASTR	Solar energy losses

[66] This RAFFELT 99 limit applies to all neutrino flavors which are light enough (<5 keV) to be emitted from globular-cluster red giants.

[67] This RAFFELT 99 limit is derived from the helioseismological limit on a new energy-loss channel of the Sun, and applies to all neutrino flavors which are light enough (<1 keV) to be emitted from the sun.

[68] BABU 94 use COOPER-SARKAR 92 limit on ν magnetic moment to derive quoted result. It applies to ν_τ.

[69] DAVIDSON 91 use data from early SLAC electron beam dump experiment to derive charge limit as a function of neutrino mass. It applies to ν_τ.

[70] Precise BARBIELLINI 87 limit depends on assumptions about the intergalactic or galactic magnetic fields and about the direct distance and time through the field. It applies to ν_e.

[71] The limit applies to all flavors.

ν (MEAN LIFE) / MASS

Measures $\left[\sum |U_{\ell j}|^2\ \Gamma_j\ m_j\right]^{-1}$, where the sum is over mass eigenstates which cannot be resolved experimentally. Some of the limits constrain the radiative decay and are based on the limit of the corresponding photon flux. Other apply to the decay of a heavier neutrino into the lighter one and a Majoron or other invisible particle. Many of these limits apply to any ν within the indicated mass range.

VALUE (s/eV)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
> 15.4	90	[72] KRAKAUER	91	CNTR	ν_μ, $\overline{\nu}_\mu$ at LAMPF
$> 7 \times 10^{9}$		[73] RAFFELT	85	ASTR	
> 300	90	[74] REINES	74	CNTR	$\overline{\nu}_e$
• • • We do not use the following data for averages, fits, limits, etc. • • •					
> 0.004	90	[75] AHARMIM	04	SNO	quasidegen. ν masses
$> 4.4 \times 10^{-5}$	90	[75] AHARMIM	04	SNO	hierarchical ν masses
$\gtrsim 100$	95	[76] CECCHINI	04	ASTR	Radiative decay for ν mass > 0.01 eV
> 0.067	90	[77] EGUCHI	04	KLND	quasidegen. ν masses
$> 1.1 \times 10^{-3}$	90	[77] EGUCHI	04	KLND	hierarchical ν masses
$> 8.7 \times 10^{-5}$	99	[78] BANDYOPA...	03	FIT	nonradiative decay
≥ 4200	90	[79] DERBIN	02B	CNTR	Solar pp and Be ν
$> 2.8 \times 10^{-5}$	99	[80] JOSHIPURA	02B	FIT	nonradiative decay
		[81] DOLGOV	99	COSM	
		[82] BILLER	98	ASTR	m_ν= 0.05–1 eV
$> 2.8 \times 10^{15}$		[83,84] BLUDMAN	92	ASTR	$m_\nu < 50$ eV
none $10^{-12} - 5 \times 10^{4}$		[85] DODELSON	92	ASTR	m_ν=1–300 keV
$< 10^{-12}$ or $> 5 \times 10^{4}$		[85] DODELSON	92	ASTR	m_ν=1–300 keV
		[86] GRANEK	91	COSM	Decaying L^0
> 6.4	90	[87] KRAKAUER	91	CNTR	ν_e at LAMPF
$> 1.1 \times 10^{15}$		[88] WALKER	90	ASTR	m_ν= 0.03 – $\sim$ 2 MeV
$> 6.3 \times 10^{15}$		[84,89] CHUPP	89	ASTR	$m_\nu < 20$ eV
$> 1.7 \times 10^{15}$		[84] KOLB	89	ASTR	$m_\nu < 20$ eV
		[90] RAFFELT	89	RVUE	$\overline{\nu}$ (Dirac, Majorana)
		[91] RAFFELT	89B	ASTR	
$> 8.3 \times 10^{14}$		[92] VONFEILIT...	88	ASTR	
> 22	68	[93] OBERAUER	87		$\overline{\nu}_R$ (Dirac)
> 38	68	[93] OBERAUER	87		$\overline{\nu}$ (Majorana)
> 59	68	[93] OBERAUER	87		$\overline{\nu}_L$ (Dirac)
> 30	68	KETOV	86	CNTR	$\overline{\nu}$ (Dirac)
> 20	68	KETOV	86	CNTR	$\overline{\nu}$ (Majorana)

		[94] BINETRUY	84	COSM	$m_\nu \sim$ 1 MeV
$>$ 0.11	90	[95] FRANK	81	CNTR	$\nu\overline{\nu}$ LAMPF
$>$ 2 $\times 10^{21}$		[96] STECKER	80	ASTR	m_ν= 10–100 eV
$>$ 1.0 $\times 10^{-2}$	90	[95] BLIETSCHAU	78	HLBC	ν_μ, CERN GGM
$>$ 1.7 $\times 10^{-2}$	90	[95] BLIETSCHAU	78	HLBC	$\overline{\nu}_\mu$, CERN GGM
$<$ 3 $\times 10^{-11}$		[97] FALK	78	ASTR	m_ν <10 MeV
$>$ 2.2 $\times 10^{-3}$	90	[95] BARNES	77	DBC	ν, ANL 12-ft
		[98] COWSIK	77	ASTR	
$>$ 3. $\times 10^{-3}$	90	[95] BELLOTTI	76	HLBC	ν, CERN GGM
$>$ 1.3 $\times 10^{-2}$	90	[95] BELLOTTI	76	HLBC	$\overline{\nu}$, CERN GGM

[72] KRAKAUER 91 quotes the limit $\tau/m_{\nu_1} > (0.75a^2 + 21.65a + 26.3)$ s/eV, where a is a parameter describing the asymmetry in the neutrino decay defined as $dN_\gamma/d\cos\theta = (1/2)(1 + a\cos\theta)$ The parameter $a = 0$ for a Majorana neutrino, but can vary from -1 to 1 for a Dirac neutrino. The bound given by the authors is the most conservative (which applies for $a = -1$).

[73] RAFFELT 85 limit on the radiative decay is from solar x- and γ-ray fluxes. Limit depends on ν flux from pp, now established from GALLEX and SAGE to be $>$ 0.5 of expectation.

[74] REINES 74 looked for ν of nonzero mass decaying radiatively to a neutral of lesser mass $+ \gamma$. Used liquid scintillator detector near fission reactor. Finds lab lifetime 6×10^7 s or more. Above value of (mean life)/mass assumes average effective neutrino energy of 0.2 MeV. To obtain the limit 6×10^7 s REINES 74 assumed that the full $\overline{\nu}_e$ reactor flux could be responsible for yielding decays with photon energies in the interval 0.1 MeV – 0.5 MeV. This represents some overestimate so their lower limit is an over-estimate of the lab lifetime (VOGEL 84). If so, OBERAUER 87 may be comparable or better.

[75] AHARMIM 04 obtained these results from the solar $\overline{\nu}_e$ flux limit set by the SNO measurement assuming ν_2 decay through nonradiative process $\nu_2 \rightarrow \overline{\nu}_1 X$, where X is a Majoron or other invisible particle. Limits are given for the cases of quasidegenerate and hierarchical neutrino masses.

[76] CECCHINI 04 obtained this bound through the observations performed on the occasion of the 21 June 2001 total solar eclipse, looking for visible photons from radiative decays of solar neutrinos. Limit is a τ/m_{ν_2} in $\nu_2 \rightarrow \nu_1\gamma$. Limit ranges from $\sim$ 100 to 10^7 s/eV for $0.01 < m_{\nu_1} < 0.1$ eV.

[77] EGUCHI 04 obtained these results from the solar $\overline{\nu}_e$ flux limit set by the KamLAND measurement assuming ν_2 decay through nonradiative process $\nu_2 \rightarrow \overline{\nu}_1 X$, where X is a Majoron or other invisible particle. Limits are given for the cases of quasidegenerate and hierarchical neutrino masses.

[78] The ratio of the lifetime over the mass derived by BANDYOPADHYAY 03 is for ν_2. They obtained this result using the following solar-neutrino data: total rates measured in Cl and Ga experiments, the Super-Kamiokande's zenith-angle spectra, and SNO's day and night spectra. They assumed that ν_1 is the lowest mass, stable or nearly stable neutrino state and ν_2 decays through nonradiative Majoron emission process, $\nu_2 \rightarrow \overline{\nu}_1 + J$, or through nonradiative process with all the final state particles being sterile. The best fit is obtained in the region of the LMA solution.

[79] DERBIN 02B (also BACK 03B) obtained this bound for the radiative decay from the results of background measurements with Counting Test Facility (the prototype of the Borexino detector). The laboratory gamma spectrum is given as $dN_\gamma/d\cos\theta = (1/2)(1 + \alpha\cos\theta)$ with α=0 for a Majorana neutrino, and α varying to -1 to 1 for a Dirac neutrino. The listed bound is for the case of α=0. The most conservative bound 1.5×10^3 s eV^{-1} is obtained for the case of $\alpha = -1$.

[80] The ratio of the lifetime over the mass derived by JOSHIPURA 02B is for ν_2. They obtained this result from the total rates measured in all solar neutrino experiments.

They assumed that ν_1 is the lowest mass, stable or nearly stable neutrino state and ν_2 decays through nonradiative process like Majoron emission decay, $\nu_2 \rightarrow \nu_1' + J$ where ν_1' state is sterile. The exact limit depends on the specific solution of the solar neutrino problem. The quoted limit is for the LMA solution.

[81] DOLGOV 99 places limits in the (Majorana) τ-associated ν mass-lifetime plane based on nucleosynthesis. Results would be considerably modified if neutrino oscillations exist.

[82] BILLER 98 use the observed TeV γ-ray spectra to set limits on the mean life of any radiatively decaying neutrino between 0.05 and 1 eV. Curve shows $\tau_\nu/\mathrm{B}_\gamma > 0.15\times10^{21}$ s at 0.05 eV, $> 1.2\times10^{21}$ s at 0.17 eV, $> 3\times10^{21}$ s at 1 eV, where B_γ is the branching ratio to photons.

[83] BLUDMAN 92 sets additional limits by this method for higher mass ranges. Cosmological limits are also obtained.

[84] Limit on the radiative decay based on nonobservation of γ's in coincidence with ν's from SN 1987A.

[85] DODELSON 92 range is for wrong-helicity keV mass Dirac ν's from the core of neutron star in SN 1987A decaying to ν's that would have interacted in KAM2 or IMB detectors.

[86] GRANEK 91 considers heavy neutrino decays to $\gamma\nu_L$ and $3\nu_L$, where $m_{\nu_L} <$100 keV. Lifetime is calculated as a function of heavy neutrino mass, branching ratio into $\gamma\nu_L$, and m_{ν_L}.

[87] KRAKAUER 91 quotes the limit for ν_e, $\tau/m_\nu > (0.3a^2 + 9.8a + 15.9)$ s/eV, where a is a parameter describing the asymmetry in the radiative neutrino decay defined as $dN_\gamma/d\cos\theta = (1/2)(1 + a\cos\theta)$ $a=0$ for a Majorana neutrino, but can vary from -1 to 1 for a Dirac neutrino. The bound given by the authors is the most conservative (which applies for $a = -1$).

[88] WALKER 90 uses SN 1987A γ flux limits after 289 days.

[89] CHUPP 89 should be multiplied by a branching ratio (about 1) and a detection efficiency (about 1/4), and pertains to radiative decay of any neutrino to a lighter or sterile neutrino.

[90] RAFFELT 89 uses KYULDJIEV 84 to obtain $\tau m^3 > 3\times10^{18}$ s eV^3 (based on $\overline{\nu}_e e^-$ cross sections). The bound for the radiative decay is not valid if electric and magnetic transition moments are equal for Dirac neutrinos.

[91] RAFFELT 89B analyze stellar evolution and exclude the region $3\times10^{12} < \tau m^3 < 3\times10^{21}$ s eV^3.

[92] Model-dependent theoretical analysis of SN 1987A neutrinos. Quoted limit is for $\left[\sum_j |U_{\ell j}|^2\, \Gamma_j\, m_j\right]^{-1}$, where $\ell=\mu$, τ. Limit is 3.3×10^{14} s/eV for $\ell=e$.

[93] OBERAUER 87 looks for photons and e^+e^- pairs from radiative decays of reactor neutrinos.

[94] BINETRUY 84 finds $\tau < 10^8$ s for neutrinos in a radiation-dominated universe.

[95] These experiments look for $\nu_k \rightarrow \nu_j\gamma$ or $\overline{\nu}_k \rightarrow \overline{\nu}_j\gamma$.

[96] STECKER 80 limit based on UV background; result given is $\tau > 4\times10^{22}$ s at m_ν=20 eV.

[97] FALK 78 finds lifetime constraints based on supernova energetics.

[98] COWSIK 77 considers variety of scenarios. For neutrinos produced in the big bang, present limits on optical photon flux require $\tau > 10^{23}$ s for $m_\nu \sim 1$ eV. See also COWSIK 79 and GOLDMAN 79.

ν MAGNETIC MOMENT

The coupling of neutrinos to an electromagnetic field is a characterized by a 3×3 matrix λ of the magnetic (μ) and electric (d) dipole moments ($\lambda = \mu$ - id). For Majorana neutrinos the matrix λ is antisymmetric and only transition moments are allowed, while for Dirac neutrinos λ is a general 3×3 matrix. In the standard electroweak theory extended to include neutrino masses (see Fujikawa 80) $\mu_\nu = 3eG_F m_\nu/(8\pi^2\sqrt{2}) = 3.2 \times 10^{-19}(m_\nu/\text{eV})\mu_B$, i.e. it is unobservably small given the known small neutrino masses. In more general models there is no longer a proportionality between neutrino mass and its magnetic moment, even though only massive neutrinos have nonvanishing magnetic moments without fine tuning.

Laboratory bounds on λ are obtained via elastic $\nu - e$ scattering, where the scattered neutrino is not observed. The combinations of matrix elements of λ that are constrained by various experiments depend on the initial neutrino flavor and on its propagation between source and detector (e.g., solar ν_e and reactor $\overline{\nu}_e$ do not constrain the same combinations). The listings below therefore identify the initial neutrino flavor.

Other limits, e.g. from various stellar cooling processes, apply to all neutrino flavors. Analogous flavor independent, but weaker, limits are obtained from the analysis of $e^+e^- \to \nu\overline{\nu}\gamma$ collider experiments.

VALUE (10^{-10} μ_B)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
< 0.9 (CL = 90%) OUR LIMIT					
< 0.9	90	[99] DARAKTCH...	05		Reactor $\overline{\nu}_e$
< 6.8	90	[100] AUERBACH	01	LSND	ν_{ee}, $\nu_{\mu e}$ scattering
< 3900	90	[101] SCHWIENHO...	01	DONU	$\nu_\tau e^- \to \nu_\tau e^-$
• • • We do not use the following data for averages, fits, limits, etc. • • •					
< 37	95	[102] GRIFOLS	04	FIT	Solar ^{8}B ν (SNO NC)
< 3.6	90	[103] LIU	04	SKAM	Solar ν spectrum shape
< 1.1	90	[104] LIU	04	SKAM	Solar ν spectrum shape (LMA region)
< 5.5	90	[105] BACK	03B	CNTR	Solar pp and Be ν
< 1.0	90	[106] DARAKTCH...	03		Reactor $\overline{\nu}_e$
< 1.3	90	[107] LI	03B	CNTR	Reactor $\overline{\nu}_e$
< 2	90	[108] GRIMUS	02	FIT	solar + reactor (Majorana ν)
<80000	90	[109] TANIMOTO	00	RVUE	$e^+e^- \to \nu\overline{\nu}\gamma$
< 0.01–0.04		[110] AYALA	99	ASTR	$\nu_L \to \nu_R$ in SN 1987A
< 1.5	90	[111] BEACOM	99	SKAM	ν spectrum shape
< 0.03		[112] RAFFELT	99	ASTR	Red giant luminosity
< 4		[113] RAFFELT	99	ASTR	Solar cooling
<44000	90	ABREU	97J	DLPH	$e^+e^- \to \nu\overline{\nu}\gamma$ at LEP
<33000	90	[114] ACCIARRI	97Q	L3	$e^+e^- \to \nu\overline{\nu}\gamma$ at LEP
< 0.62		[115] ELMFORS	97	COSM	Depolarization in early universe plasma
<27000	95	[116] ESCRIBANO	97	RVUE	$\Gamma(Z \to \nu\nu)$ at LEP

<30	90	VILAIN	95B	CHM2	$\nu_\mu e \to \nu_\mu e$
<55000	90	GOULD	94	RVUE	$e^+ e^- \to \nu \overline{\nu} \gamma$ at LEP
<1.9	95	[117] DERBIN	93	CNTR	Reactor $\overline{\nu} e \to \overline{\nu} e$
<5400	90	[118] COOPER-...	92	BEBC	$\nu_\tau e^- \to \nu_\tau e^-$
<2.4	90	[119] VIDYAKIN	92	CNTR	Reactor $\overline{\nu} e \to \overline{\nu} e$
<56000	90	DESHPANDE	91	RVUE	$e^+ e^- \to \nu \overline{\nu} \gamma$
<100	95	[120] DORENBOS...	91	CHRM	$\nu_\mu e \to \nu_\mu e$
<8.5	90	AHRENS	90	CNTR	$\nu_\mu e \to \nu_\mu e$
<10.8	90	[121] KRAKAUER	90	CNTR	LAMPF $\nu e \to \nu e$
<7.4	90	[121] KRAKAUER	90	CNTR	LAMPF $(\nu_\mu, \overline{\nu}_\mu) e$ elast.
<0.02		[122] RAFFELT	90	ASTR	Red giant luminosity
<0.1		[123] RAFFELT	89B	ASTR	Cooling helium stars
		[124] FUKUGITA	88	COSM	Primordial magn. fields
<40000	90	[125] GROTCH	88	RVUE	$e^+ e^- \to \nu \overline{\nu} \gamma$
$\leq .3$		[123] RAFFELT	88B	ASTR	He burning stars
<0.11		[123] FUKUGITA	87	ASTR	Cooling helium stars
<0.0006		[126] NUSSINOV	87	ASTR	Cosmic EM backgrounds
<0.1–0.2		MORGAN	81	COSM	^{4}He abundance
<0.85		BEG	78	ASTR	Stellar plasmons
<0.6		[127] SUTHERLAND	76	ASTR	Red giants + degenerate dwarfs
<81		[128] KIM	74	RVUE	$\overline{\nu}_\mu e \to \overline{\nu}_\mu e$
<1		BERNSTEIN	63	ASTR	Solar cooling
<14		COWAN	57	CNTR	Reactor $\overline{\nu}$

[99] DARAKTCHIEVA 05 present the final analysis of the search for non-standard $\overline{\nu}_e - e$ scattering component at Bugey nuclear reactor. Full kinematical event reconstruction of both the kinetic energy above 700 keV and scattering angle of the recoil electron, by use of TPC. Most stringent laboratory limit on magnetic moment. Supersedes DARAKTCHIEVA 03.

[100] AUERBACH 01 limit is based on the LSND ν_e and ν_μ electron scattering measurements. The limit is slightly more stringent than KRAKAUER 90.

[101] SCHWIENHORST 01 quote an experimental sensitivity of 4.9×10^{-7}.

[102] GRIFOLS 04 obtained this bound using the SNO data of the solar ^{8}B neutrino flux measured with deuteron breakup. This bound applies to $\mu_{\text{eff}} = (\mu_{21}^2 + \mu_{22}^2 + \mu_{23}^2)^{1/2}$.

[103] LIU 04 obtained this limit using the shape of the recoil electron energy spectrum from the Super-Kamiokande-I 1496 days of solar neutrino data. Neutrinos are assumed to have only diagonal magnetic moments, $\mu_{\nu 1} = \mu_{\nu 2}$. This limit corresponds to the oscillation parameters in the vacuum oscillation region.

[104] LIU 04 obtained this limit using the shape of the recoil electron energy spectrum from the Super-Kamiokande-I 1496 live-day solar neutrino data, by limiting the oscillation parameter region in the LMA region allowed by solar neutrino experiments plus KamLAND. $\mu_{\nu 1} = \mu_{\nu 2}$ is assumed. In the LMA region, the same limit would be obtained even if neutrinos have off-diagonal magnetic moments.

[105] BACK 03B obtained this bound from the results of background measurements with Counting Test Facility (the prototype of the Borexino detector). Standard Solar Model flux was assumed. This μ_ν can be different from the reactor μ_ν in certain oscillation scenarios (see BEACOM 99).

[106] DARAKTCHIEVA 03 searched for non-standard $\overline{\nu}_e$-e scattering component at Bugey nuclear reactor. Full kinematical event reconstruction by use of TPC. Superseded by DARAKTCHIEVA 05.

[107] LI 03B used Ge detector in active shield near nuclear reactor to test for nonstandard $\overline{\nu}_e$-e scattering.

[108] GRIMUS 02 obtain stringent bounds on all Majorana neutrino transition moments from a simultaneous fit of LMA-MSW oscillation parameters and transition moments to global solar neutrino data + reactor data. Using only solar neutrino data, a 90% CL bound of $6.3 \times 10^{-10} \mu_B$ is obtained.

[109] TANIMOTO 00 combined $e^+ e^- \rightarrow \nu \overline{\nu} \gamma$ data from VENUS, TOPAZ, and AMY.

[110] AYALA 99 improves the limit of BARBIERI 88.

[111] BEACOM 99 obtain the limit using the shape, but not the absolute magnitude which is affected by oscillations, of the solar neutrino spectrum obtained by Superkamiokande (825 days). This μ_ν can be different from the reactor μ_ν in certain oscillation scenarios.

[112] RAFFELT 99 is an update of RAFFELT 90. This limit applies to all neutrino flavors which are light enough (< 5 keV) to be emitted from globular-cluster red giants. This limit pertains equally to electric dipole moments and magnetic transition moments, and it applies to both Dirac and Majorana neutrinos.

[113] RAFFELT 99 is essentially an update of BERNSTEIN 63, but is derived from the helioseismological limit on a new energy-loss channel of the Sun. This limit applies to all neutrino flavors which are light enough (<1 keV) to be emitted from the Sun. This limit pertains equally to electric dipole and magnetic transition moments, and it applies to both Dirac and Majorana neutrinos.

[114] ACCIARRI 97Q result applies to both direct and transition magnetic moments and for $q^2=0$.

[115] ELMFORS 97 calculate the rate of depolarization in a plasma for neutrinos with a magnetic moment and use the constraints from a big-bang nucleosynthesis on additional degrees of freedom.

[116] Applies to absolute value of magnetic moment.

[117] DERBIN 93 determine the cross section for 0.6–2.0 MeV electron energy as $(1.28 \pm 0.63) \times \sigma_{\text{weak}}$. However, the (reactor on − reactor off)/(reactor off) is only $\sim 1/100$.

[118] COOPER-SARKAR 92 assume $f_{D_s}/f_\pi = 2$ and D_s, $\overline{D}_s$ production cross section = 2.6 μb to calculate ν flux.

[119] VIDYAKIN 92 limit is from a $e\overline{\nu}_e$ elastic scattering experiment. No experimental details are given except for the cross section from which this limit is derived. Signal/noise was 1/10. The limit uses $\sin^2\theta_W = 0.23$ as input.

[120] DORENBOSCH 91 corrects an incorrect statement in DORENBOSCH 89 that the ν magnetic moment is $< 1 \times 10^{-9}$ at the 95%CL. DORENBOSCH 89 measures both $\nu_\mu e$ and $\overline{\nu} e$ elastic scattering and assume $\mu(\nu) = \mu(\overline{\nu})$.

[121] KRAKAUER 90 experiment fully reported in ALLEN 93.

[122] RAFFELT 90 limit applies for a diagonal magnetic moment of a Dirac neutrino, or for a transition magnetic moment of a Majorana neutrino. In the latter case, the same analysis gives $< 1.4 \times 10^{-12}$. Limit at 95%CL obtained from δM_c.

[123] Significant dependence on details of stellar models.

[124] FUKUGITA 88 find magnetic dipole moments of any two neutrino species are bounded by $\mu < 10^{-16}$ $[10^{-9}$ $G/B_0]$ where B_0 is the present-day intergalactic field strength.

[125] GROTCH 88 combined data from MAC, ASP, CELLO, and Mark J.

[126] For $m_\nu = 8$–200 eV. NUSSINOV 87 examines transition magnetic moments for $\nu_\mu \rightarrow \nu_e$ and obtain $< 3 \times 10^{-15}$ for $m_\nu > 16$ eV and $< 6 \times 10^{-14}$ for $m_\nu > 4$ eV.

[127] We obtain above limit from SUTHERLAND 76 using their limit $f < 1/3$.

[128] KIM 74 is a theoretical analysis of $\overline{\nu}_\mu$ reaction data.

NEUTRINO CHARGE RADIUS SQUARED

We report limits on the so-called neutrino charge radius squared. While the straight-forward definition of a neutrino charge radius has been proven to be gauge-dependent and, hence, unphysical (LEE 77C), there have been recent attempts to define a physically observable neutrino charge radius (BERNABEU 00, BERNABEU 02). The issue is still controversial (FUJIKAWA 03, BERNABEU 03). A more general interpretation of the experimental results is that they are limits on certain nonstandard contributions to neutrino scattering.

VALUE (10^{-32} cm^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
-2.97 to 4.14	90	[129] AUERBACH	01	LSND	$\nu_e e \to \nu_e e$

• • • We do not use the following data for averages, fits, limits, etc. • • •

VALUE (10^{-32} cm^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
<0.68, > -0.53	90	[130] HIRSCH	03		$\nu_\mu e$ scat.
<9.9 and > -8.2	90	[131] HIRSCH	03		anomalous $e^+ e^- \to \nu \overline{\nu} \gamma$
$< \lvert 0.6 \rvert$	90	VILAIN	95B	CHM2	$\nu_\mu e$ elastic scat.
0.9 ± 2.7		ALLEN	93	CNTR	LAMPF $\nu e \to \nu e$
< 2.3	95	MOURAO	92	ASTR	HOME/KAM2 ν rates
< 7.3	90	[132] VIDYAKIN	92	CNTR	Reactor $\overline{\nu} e \to \overline{\nu} e$
1.1 ± 2.3		ALLEN	91	CNTR	Repl. by ALLEN 93
-1.1 ± 1.0		[133] AHRENS	90	CNTR	$\nu_\mu e$ elastic scat.
-0.3 ± 1.5		[133] DORENBOS...	89	CHRM	$\nu_\mu e$ elastic scat.
		[134] GRIFOLS	89B	ASTR	SN 1987A

[129] AUERBACH 01 measure $\nu_e e$ elastic scattering with LSND detector. The cross section agrees with the Standard Model expectation, including the charge and neutral current interference. The 90% CL applies to the range shown.

[130] Based on analysis of CCFR 98 results. Limit is on $\langle r_V^2 \rangle + \langle r_A^2 \rangle$. The CHARM II and E734 at BNL results are reanalyzed, and weaker bounds on the charge radius squared than previously published are obtained. The NuTeV result is discussed; when tentatively interpreted as ν_μ charge radius it implies $\langle r_V^2 \rangle + \langle r_A^2 \rangle = (4.20 \pm 1.64) \times 10^{-33}$ cm^2.

[131] Results of LEP-2 are interpreted as limits on the axial-vector charge radius squared of a Majorana ν_τ. Slightly weaker limits for both vector and axial-vector charge radius squared are obtained for the Dirac case, and somewhat weaker limits are obtained from the analysis of lower energy data (LEP-1.5 and TRISTAN).

[132] VIDYAKIN 92 limit is from a $e\overline{\nu}$ elastic scattering experiment. No experimental details are given except for the cross section from which this limit is derived. Signal/noise was 1/10. The limit uses $\sin^2\theta_W = 0.23$ as input.

[133] Result is obtained from reanalysis given in ALLEN 91, followed by our reduction to obtain $1\,\sigma$ errors.

[134] GRIFOLS 89B sets a limit of $\langle r^2 \rangle < 0.2 \times 10^{-32}$ cm^2 for right-handed neutrinos.

REFERENCES FOR Neutrino Properties

DARAKTCH...	05	PL B615 153	Z. Daraktchieva *et al.*	(MUNU Collab.)
ICHIKAWA	05	PR D71 043001	K. Ichikawa, M. Fukugita, M. Kawasaki	(ICRR)
KRAUS	05	EPJ C40 447	Ch. Kraus *et al.*	
AHARMIM	04	PR D70 093014	B. Aharmim *et al.*	(SNO Collab.)
BARGER	04	PL B595 55	V. Barger, D. Marfatia, A. Tregre	
CECCHINI	04	ASP 21 183	S. Cecchini *et al.*	(BGNA+)
CROTTY	04	PR D69 123007	P. Crotty, J. Lesgourgues, S. Pastor	
EGUCHI	04	PRL 92 071301	K. Eguchi *et al.*	(KamLAND Collab.)
GRIFOLS	04	PL B587 184	J.A. Grifols, E. Masso, S. Mohanty	(BARC, AHMED)
LIU	04	PRL 93 021802	D.W. Liu *et al.*	(Super-Kamiokande Collab.)
ARNABOLDI	03A	PRL 91 161802	C. Arnaboldi *et al.*	
BACK	03B	PL B563 35	H.O. Back *et al.*	(Borexino Collab.)
BANDYOPA...	03	PL B555 33	A. Bandyopadhyay, S. Choubey, S. Goswami	(SAHA+)
BERNABEU	03	hep-ph/0303202	J. Bernabeu, J. Papavassiliou, J. Vidal	
DARAKTCH...	03	PL B564 190	Z. Daraktchieva *et al.*	(MUNU Collab.)
FUJIKAWA	03	hep-ph/0303188	K. Fujikawa, R. Shrock	
HANNESTAD	03B	JCAP 0305 004	S. Hannestad	
HIRSCH	03	PR D67 033005	M. Hirsch *et al.*	
LI	03B	PRL 90 131802	H.B. Li *et al.*	(TEXONO Collab.)
SPERGEL	03	APJS 148 175	D.N. Spergel *et al.*	
BERNABEU	02	PRL 89 101802	J. Bernabeu, J. Papavassiliou, J. Vidal	
Also		PRL 89 229902 (erratum)	J. Bernabeu, J. Papavassiliou, J. Vidal	
DERBIN	02B	JETPL 76 409	A.V. Derbin, O.Ju. Smirnov	
		Translated from ZETFP 76 483.		
ELGAROY	02	PRL 89 061301	O. Elgaroy *et al.*	
GRIMUS	02	NP B648 376	W. Grimus *et al.*	
JOSHIPURA	02B	PR D66 113008	A.S. Joshipura, E. Masso, S. Mohanty	
LEWIS	02	PR D66 103511	A. Lewis, S. Bridle	
LOREDO	02	PR D65 063002	T.J. Loredo, D.Q. Lamb	
WANG	02	PR D65 123001	X. Wang, M. Tegmark, M. Zaldarriaga	
AUERBACH	01	PR D63 112001	L.B. Auerbach *et al.*	(LSND Collab.)
SCHWIENHO...	01	PL B513 23	R. Schwienhorst *et al.*	(DONUT Collab.)
ATHANAS	00	PR D61 052002	M. Athanas *et al.*	(CLEO Collab.)
BERNABEU	00	PR D62 113012	J. Bernabeu *et al.*	
FUKUGITA	00	PRL 84 1082	M. Fukugita, G.C. Liu, N. Sugiyama	
TANIMOTO	00	PL B478 1	N. Tanimoto *et al.*	
AYALA	99	PR D59 111901	A. Ayala, J.C. D'Olivo, M. Torres	
BEACOM	99	PRL 83 5222	J.F. Beacom, P. Vogel	
CROFT	99	PRL 83 1092	R.A.C. Croft, W. Hu, R. Dave	
DOLGOV	99	NP B548 385	A.D. Dolgov *et al.*	
LOBASHEV	99	PL B460 227	V.M. Lobashev *et al.*	
RAFFELT	99	PRPL 320 319	G.G. Raffelt	
WEINHEIMER	99	PL B460 219	Ch. Weinheimer *et al.*	
ACKERSTAFF	98T	EPJ C5 229	K. Ackerstaff *et al.*	(OPAL Collab.)
AMMAR	98	PL B431 209	R. Ammar *et al.*	(CLEO Collab.)
BARATE	98F	EPJ C2 395	R. Barate *et al.*	(ALEPH Collab.)
BILLER	98	PRL 80 2992	S.D. Biller *et al.*	(WHIPPLE Collab.)
FELDMAN	98	PR D57 3873	G.J. Feldman, R.D. Cousins	
LENZ	98	PL B416 50	S. Lenz *et al.*	
ABREU	97J	ZPHY C74 577	P. Abreu *et al.*	(DELPHI Collab.)
ACCIARRI	97Q	PL B412 201	M. Acciarri *et al.*	(L3 Collab.)
ANASTASSOV	97	PR D55 2559	A. Anastassov *et al.*	(CLEO Collab.)
Also		PR D58 119903 (erratum)	A. Anastassov *et al.*	(CLEO Collab.)
ELMFORS	97	NP B503 3	P. Elmfors *et al.*	
ESCRIBANO	97	PL B395 369	R. Escribano, E. Masso	(BARC, PARIT)
FIELDS	97	ASP 6 169	B.D. Fields, K. Kainulainen, K.A. Olive	(NDAM+)
SWAIN	97	PR D55 R1	J. Swain, L. Taylor	(NEAS)
ALEXANDER	96M	ZPHY C72 231	G. Alexander *et al.*	(OPAL Collab.)
ASSAMAGAN	96	PR D53 6065	K.A. Assamagan *et al.*	(PSI, ZURI, VILL+)
BAI	96	PR D53 20	J.Z. Bai *et al.*	(BES Collab.)
BOTTINO	96	PR D53 6361	A. Bottino *et al.*	
DOLGOV	96	PL B383 193	A.D. Dolgov, S. Pastor, J.W.F. Valle	(IFIC, VALE)
HANNESTAD	96	PRL 76 2848	S. Hannestad, J. Madsen	(AARH)
HANNESTAD	96B	PRL 77 5148 (erratum)	S. Hannestad, J. Madsen	(AARH)
HANNESTAD	96C	PR D54 7894	S. Hannestad, J. Madsen	(AARH)
SOBIE	96	ZPHY C70 383	R.J. Sobie, R.K. Keeler, I. Lawson	(VICT)
BELESEV	95	PL B350 263	A.I. Belesev *et al.*	(INRM, KIAE)
BUSKULIC	95H	PL B349 585	D. Buskulic *et al.*	(ALEPH Collab.)

CHING	95	IJMP A10 2841	C.R. Ching *et al.*	(CST, BEIJT, CIAE)
DOLGOV	95	PR D51 4129	A.D. Dolgov, K. Kainulainen, I.Z. Rothstein	(MICH+)
HIDDEMANN	95	JPG 21 639	K.H. Hiddemann, H. Daniel, O. Schwentker	(MUNT)
KERNAN	95	NP B437 243	P.J. Kernan, L.M. Krauss	(CASE)
SIGL	95	PR D51 1499	G. Sigl, M.S. Turner	(FNAL, EFI)
STOEFFL	95	PRL 75 3237	W. Stoeffl, D.J. Decman	(LLNL)
VILAIN	95B	PL B345 115	P. Vilain *et al.*	(CHARM II Collab.)
ASSAMAGAN	94	PL B335 231	K.A. Assamagan *et al.*	(PSI, ZURI, VILL+)
BABU	94	PL B321 140	K.S. Babu, T.M. Gould, I.Z. Rothstein	(BART+)
DODELSON	94	PR D49 5068	S. Dodelson, G. Gyuk, M.S. Turner	(FNAL, CHIC+)
GOULD	94	PL B333 545	T.M. Gould, I.Z. Rothstein	(JHU, MICH)
JECKELMANN	94	PL B335 326	B. Jeckelmann, P.F.A. Goudsmit, H.J. Leisi	(WABRN+)
KAWASAKI	94	NP B419 105	M. Kawasaki *et al.*	(OSU)
PERES	94	PR D50 513	O.L.G. Peres, V. Pleitez, R. Zukanovich Funchal	
YASUMI	94	PL B334 229	S. Yasumi *et al.*	(KEK, TSUK, KYOT+)
ALLEN	93	PR D47 11	R.C. Allen *et al.*	(UCI, LANL, ANL+)
BALEST	93	PR D47 R3671	R. Balest *et al.*	(CLEO Collab.)
CINABRO	93	PRL 70 3700	D. Cinabro *et al.*	(CLEO Collab.)
DERBIN	93	JETPL 57 768	A.V. Derbin *et al.*	(PNPI)
		Translated from ZETFP 57 755.		
DOLGOV	93	PRL 71 476	A.D. Dolgov, I.Z. Rothstein	(MICH)
ENQVIST	93	PL B301 376	K. Enqvist, H. Uibo	(NORD)
SUN	93	CJNP 15 261	H.C. Sun *et al.*	(CIAE, CST, BEIJT)
WEINHEIMER	93	PL B300 210	C. Weinheimer *et al.*	(MANZ)
ALBRECHT	92M	PL B292 221	H. Albrecht *et al.*	(ARGUS Collab.)
BLUDMAN	92	PR D45 4720	S.A. Bludman	(CFPA)
COOPER-...	92	PL B280 153	A.M. Cooper-Sarkar *et al.*	(BEBC WA66 Collab.)
DODELSON	92	PRL 68 2572	S. Dodelson, J.A. Frieman, M.S. Turner	(FNAL+)
HOLZSCHUH	92B	PL B287 381	E. Holzschuh, M. Fritschi, W. Kundig	(ZURI)
KAWANO	92	PL B275 487	L.H. Kawano *et al.*	(CIT, UCSD, LLL+)
MOURAO	92	PL B285 364	A.M. Mourao, J. Pulido, J.P. Ralston	(LISB, LISBT+)
PDG	92	PR D45, 1 June, Part II	K. Hikasa *et al.*	(KEK, LBL, BOST+)
VIDYAKIN	92	JETPL 55 206	G.S. Vidyakin *et al.*	(KIAE)
		Translated from ZETFP 55 212.		
ALLEN	91	PR D43 R1	R.C. Allen *et al.*	(UCI, LANL, UMD)
DAVIDSON	91	PR D43 2314	S. Davidson, B.A. Campbell, D. Bailey	(ALBE+)
DESHPANDE	91	PR D43 943	N.G. Deshpande, K.V.L. Sarma	(OREG, TATA)
DORENBOS...	91	ZPHY C51 142	J. Dorenbosch *et al.*	(CHARM Collab.)
FULLER	91	PR D43 3136	G.M. Fuller, R.A. Malaney	(UCSD)
GRANEK	91	IJMP A6 2387	H. Granek, B.H.J. McKellar	(MELB)
KAWAKAMI	91	PL B256 105	H. Kawakami *et al.*	(INUS, TOHOK, TINT+)
KOLB	91	PRL 67 533	E.W. Kolb *et al.*	(FNAL, CHIC)
KRAKAUER	91	PR D44 R6	D.A. Krakauer *et al.*	(LAMPF E225 Collab.)
LAM	91	PR D44 3345	W.P. Lam, K.W. Ng	(AST)
ROBERTSON	91	PRL 67 957	R.G.H. Robertson *et al.*	(LASL, LLL)
AHRENS	90	PR D41 3297	L.A. Ahrens *et al.*	(BNL, BROW, HIRO+)
AVIGNONE	90	PR D41 682	F.T. Avignone, J.I. Collar	(SCUC)
KRAKAUER	90	PL B252 177	D.A. Krakauer *et al.*	(LAMPF E225 Collab.)
RAFFELT	90	PRL 64 2856	G.G. Raffelt	(MPIM)
WALKER	90	PR D41 689	T.P. Walker	(HARV)
CHUPP	89	PRL 62 505	E.L. Chupp, W.T. Vestrand, C. Reppin	(UNH, MPIM)
DORENBOS...	89	ZPHY C41 567	J. Dorenbosch *et al.*	(CHARM Collab.)
GRIFOLS	89B	PR D40 3819	J.A. Grifols, E. Masso	(BARC)
KOLB	89	PRL 62 509	E.W. Kolb, M.S. Turner	(CHIC, FNAL)
LOREDO	89	ANYAS 571 601	T.J. Loredo, D.Q. Lamb	(CHIC)
RAFFELT	89	PR D39 2066	G.G. Raffelt	(PRIN, UCB)
RAFFELT	89B	APJ 336 61	G. Raffelt, D. Dearborn, J. Silk	(UCB, LLL)
ALBRECHT	88B	PL B202 149	H. Albrecht *et al.*	(ARGUS Collab.)
BARBIERI	88	PRL 61 27	R. Barbieri, R.N. Mohapatra	(PISA, UMD)
BORIS	88	PRL 61 245 (erratum)	S.D. Boris *et al.*	(ITEP, ASCI)
FUKUGITA	88	PRL 60 879	M. Fukugita *et al.*	(KYOTU, MPIM, UCB)
GROTCH	88	ZPHY C39 553	H. Grotch, R.W. Robinett	(PSU)
RAFFELT	88B	PR D37 549	G.G. Raffelt, D.S.P. Dearborn	(UCB, LLL)
SPERGEL	88	PL B200 366	D.N. Spergel, J.N. Bahcall	(IAS)
VONFEILIT...	88	PL B200 580	F. von Feilitzsch, L. Oberauer	(MUNT)
BARBIELLINI	87	NAT 329 21	G. Barbiellini, G. Cocconi	(CERN)
BORIS	87	PRL 58 2019	S.D. Boris *et al.*	(ITEP, ASCI)
Also		PRL 61 245 (erratum)	S.D. Boris *et al.*	(ITEP, ASCI)
BORIS	87B	JETPL 45 333	S.D. Boris *et al.*	(ITEP)
		Translated from ZETFP 45 267.		

FUKUGITA	87	PR D36 3817	M. Fukugita, S. Yazaki	(KYOTU, TOKY)
NUSSINOV	87	PR D36 2278	S. Nussinov, Y. Rephaeli	(TELA)
OBERAUER	87	PL B198 113	L.F. Oberauer, F. von Feilitzsch, R.L. Mossbauer	
SPRINGER	87	PR A35 679	P.T. Springer *et al.*	(LLNL)
KETOV	86	JETPL 44 146	S.N. Ketov *et al.*	(KIAE)
		Translated from ZETFP 44 114.		
COWSIK	85	PL 151B 62	R. Cowsik	(TATA)
RAFFELT	85	PR D31 3002	G.G. Raffelt	(MPIM)
BINETRUY	84	PL 134B 174	P. Binetruy, G. Girardi, P. Salati	(LAPP)
FREESE	84	NP B233 167	K. Freese, D.N. Schramm	(CHIC, FNAL)
KYULDJIEV	84	NP B243 387	A.V. Kyuldjiev	(SOFI)
SCHRAMM	84	PL 141B 337	D.N. Schramm, G. Steigman	(FNAL, BART)
VOGEL	84	PR D30 1505	P. Vogel	
ANDERHUB	82	PL 114B 76	H.B. Anderhub *et al.*	(ETH, SIN)
OLIVE	82	PR D25 213	K.A. Olive, M.S. Turner	(CHIC, UCSB)
BERNSTEIN	81	PL 101B 39	J. Bernstein, G. Feinberg	(STEV, COLU)
FRANK	81	PR D24 2001	J.S. Frank *et al.*	(LASL, YALE, MIT+)
MORGAN	81	PL 102B 247	J.A. Morgan	(SUSS)
LUBIMOV	80	PL 94B 266	V.A. Lyubimov *et al.*	(ITEP)
STECKER	80	PRL 45 1460	F.W. Stecker	(NASA)
COWSIK	79	PR D19 2219	R. Cowsik	(TATA)
GOLDMAN	79	PR D19 2215	T. Goldman, G.J. Stephenson	(LASL)
BEG	78	PR D17 1395	M.A.B. Beg, W.J. Marciano, M. Ruderman	(ROCK+)
BLIETSCHAU	78	NP B133 205	J. Blietschau *et al.*	(Gargamelle Collab.)
FALK	78	PL 79B 511	S.W. Falk, D.N. Schramm	(CHIC)
BARNES	77	PRL 38 1049	V.E. Barnes *et al.*	(PURD, ANL)
COWSIK	77	PRL 39 784	R. Cowsik	(MPIM, TATA)
LEE	77C	PR D16 1444	B.W. Lee, R.E. Shrock	(STON)
VYSOTSKY	77	JETPL 26 188	M.I. Vysotsky, A.D. Dolgov, Y.B. Zeldovich	(ITEP)
		Translated from ZETFP 26 200.		
BELLOTTI	76	LNC 17 553	E. Bellotti *et al.*	(MILA)
SUTHERLAND	76	PR D13 2700	P. Sutherland *et al.*	(PENN, COLU, NYU)
SZALAY	76	AA 49 437	A.S. Szalay, G. Marx	(EOTV)
CLARK	74	PR D9 533	A.R. Clark *et al.*	(LBL)
KIM	74	PR D9 3050	J.E. Kim, V.S. Mathur, S. Okubo	(ROCH)
REINES	74	PRL 32 180	F. Reines, H.W. Sobel, H.S. Gurr	(UCI)
SZALAY	74	APAH 35 8	A.S. Szalay, G. Marx	(EOTV)
COWSIK	72	PRL 29 669	R. Cowsik, J. McClelland	(UCB)
MARX	72	Nu Conf. Budapest	G. Marx, A.S. Szalay	(EOTV)
GERSHTEIN	66	JETPL 4 120	S.S. Gershtein, Y.B. Zeldovich	(KIAM)
		Translated from ZETFP 4 189.		
BERNSTEIN	63	PR 132 1227	J. Bernstein, M. Ruderman, G. Feinberg	(NYU+)
COWAN	57	PR 107 528	C.L. Cowan, F. Reines	(LANL)