

# Number of Neutrino Types

The neutrinos referred to in this section are those of the Standard  $SU(2)\times U(1)$  Electroweak Model possibly extended to allow nonzero neutrino masses. Light neutrinos are those with  $m < m_Z/2$ . The limits are on the number of neutrino mass eigenstates, including  $\nu_1$ ,  $\nu_2$ , and  $\nu_3$ .

## THE NUMBER OF LIGHT NEUTRINO TYPES FROM COLLIDER EXPERIMENTS

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The most precise measurements of the number of light neutrino types,  $N_\nu$ , come from studies of  $Z$  production in  $e^+e^-$  collisions. The invisible partial width,  $\Gamma_{\text{inv}}$ , is determined by subtracting the measured visible partial widths, corresponding to  $Z$  decays into quarks and charged leptons, from the total  $Z$  width. The invisible width is assumed to be due to  $N_\nu$  light neutrino species each contributing the neutrino partial width  $\Gamma_\nu$  as given by the Standard Model. In order to reduce the model dependence, the Standard Model value for the ratio of the neutrino to charged leptonic partial widths,  $(\Gamma_\nu/\Gamma_\ell)_{\text{SM}} = 1.991 \pm 0.001$ , is used instead of  $(\Gamma_\nu)_{\text{SM}}$  to determine the number of light neutrino types:

$$N_\nu = \frac{\Gamma_{\text{inv}}}{\Gamma_\ell} \left( \frac{\Gamma_\ell}{\Gamma_\nu} \right)_{\text{SM}} . \quad (1)$$

The combined result from the four LEP experiments is  $N_\nu = 2.984 \pm 0.008$  [1]. Recent analyses applied corrections to the LEP result [1] by including the effect of correlated luminosity systematics and also using an improved Bhabha cross section calculation [2,3] to obtain  $N_\nu = 2.9963 \pm 0.0074$ .

In the past, when only small samples of  $Z$  decays had been recorded by the LEP experiments and by the Mark II at SLC,

the uncertainty in  $N_\nu$  was reduced by using Standard Model fits to the measured hadronic cross sections at several center-of-mass energies near the  $Z$  resonance. Since this method is much more dependent on the Standard Model, the approach described above is favored.

Before SLC and LEP, limits on the number of neutrino generations were placed by experiments at lower-energy  $e^+e^-$  colliders by measuring the cross section of the process  $e^+e^- \rightarrow \nu\bar{\nu}\gamma$ . The ASP, CELLO, MAC, MARK J, and VENUS experiments observed a total of 3.9 events above background [4], leading to a 95% CL limit of  $N_\nu < 4.8$ . This process has a much larger cross section at center-of-mass energies near the  $Z$  mass and has been measured at LEP by the ALEPH, DELPHI, L3, and OPAL experiments [5]. These experiments have observed several thousand such events, and the combined result is  $N_\nu = 3.00 \pm 0.08$ . The same process has also been measured by the LEP experiments at much higher center-of-mass energies, between 130 and 208 GeV, in searches for new physics [6]. Combined with the lower energy data, the result is  $N_\nu = 2.92 \pm 0.05$ .

Experiments at  $p\bar{p}$  colliders also placed limits on  $N_\nu$  by determining the total  $Z$  width from the observed ratio of  $W^\pm \rightarrow \ell^\pm\nu$  to  $Z \rightarrow \ell^+\ell^-$  events [7]. This involved a calculation that assumed Standard Model values for the total  $W$  width and the ratio of  $W$  and  $Z$  leptonic partial widths, and used an estimate of the ratio of  $Z$  to  $W$  production cross sections. Now that the  $Z$  width is very precisely known from the LEP experiments, the approach is now one of those used to determine the  $W$  width.

## References

1. ALEPH, DELPHI, L3, OPAL, and SLD Collaborations, and LEP Electroweak Working Group, and SLD Electroweak

- Group, and SLD Heavy Flavour Group, Phys. Reports **427**, 257 (2006).
2. P. Janot and S. Jadach, Phys. Lett. **B803**, 135319 (2020).
  3. G. Voutsinas *et al.*, Phys. Lett. **B800**, 135068 (2020).
  4. VENUS: K. Abe *et al.*, Phys. Lett. **B232**, 431 (1989);  
 ASP: C. Hearty *et al.*, Phys. Rev. **D39**, 3207 (1989);  
 CELLO: H.J. Behrend *et al.*, Phys. Lett. **B215**, 186 (1988);  
 MAC: W.T. Ford *et al.*, Phys. Rev. **D33**, 3472 (1986);  
 MARK J: H. Wu, Ph.D. Thesis, Univ. Hamburg (1986).
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 DELPHI: P. Abreu *et al.*, Z. Phys. **C74**, 577 (1997);  
 OPAL: R. Akers *et al.*, Z. Phys. **C65**, 47 (1995);  
 ALEPH: D. Buskulic *et al.*, Phys. Lett. **B313**, 520 (1993).
  6. DELPHI: J. Abdallah *et al.*, Eur. Phys. J. **C38**, 395 (2005);  
 L3: P. Achard *et al.*, Phys. Lett. **B587**, 16 (2004);  
 ALEPH: A. Heister *et al.*, Eur. Phys. J. **C28**, 1 (2003);  
 OPAL: G. Abbiendi *et al.*, Eur. Phys. J. **C18**, 253 (2000).
  7. UA1: C. Albajar *et al.*, Phys. Lett. **B198**, 271 (1987);  
 UA2: R. Ansari *et al.*, Phys. Lett. **B186**, 440 (1987).

### Number from $e^+e^-$ Colliders

#### Number of Light $\nu$ Types

<u>VALUE</u>	<u>DOCUMENT ID</u>	<u>TECN</u>
<b>2.9963±0.0074</b>	<sup>1</sup> JANOT	20
• • • We do not use the following data for averages, fits, limits, etc. • • •		
2.9918±0.0081	<sup>2</sup> VOUTSINAS	20
2.9840±0.0082	<sup>3</sup> LEP-SLC	06 RVUE
3.00 ±0.05	<sup>4</sup> LEP	92 RVUE

<sup>1</sup> JANOT 20 applies a correction to LEP-SLC 06 using an updated Bhabha cross section calculation. This result also includes a correction to account for correlated luminosity bias as presented in VOUTSINAS 20.

<sup>2</sup> VOUTSINAS 20 applies a correction to LEP-SLC 06 to account for correlated luminosity bias.

<sup>3</sup> Combined fit from ALEPH, DELPHI, L3 and OPAL Experiments.

<sup>4</sup> Simultaneous fits to all measured cross section data from all four LEP experiments.

### Number of Light $\nu$ Types from Direct Measurement of Invisible Z Width

In the following, the invisible Z width is obtained from studies of single-photon events from the reaction  $e^+e^- \rightarrow \nu\bar{\nu}\gamma$ . All are obtained from LEP runs in the  $E_{\text{cm}}^{ee}$  range 88–209 GeV.

<u>VALUE</u>	<u>DOCUMENT ID</u>	<u>TECN</u>	<u>COMMENT</u>
<b>2.92±0.05 OUR AVERAGE</b>	Error includes scale factor of 1.2.		
2.84±0.10±0.14	ABDALLAH	05B DLPH	$\sqrt{s} = 180\text{--}209$ GeV
2.98±0.05±0.04	ACHARD	04E L3	1990–2000 LEP runs
2.86±0.09	HEISTER	03C ALEP	$\sqrt{s} = 189\text{--}209$ GeV
2.69±0.13±0.11	ABBIENDI,G	00D OPAL	1998 LEP run
2.89±0.32±0.19	ABREU	97J DLPH	1993–1994 LEP runs
3.23±0.16±0.10	AKERS	95C OPAL	1990–1992 LEP runs
2.68±0.20±0.20	BUSKULIC	93L ALEP	1990–1991 LEP runs
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●			
2.84±0.15±0.14	ABREU	00Z DLPH	1997–1998 LEP runs
3.01±0.08	ACCIARRI	99R L3	1991–1998 LEP runs
3.1 ±0.6 ±0.1	ADAM	96C DLPH	$\sqrt{s} = 130, 136$ GeV

### Limits from Astrophysics and Cosmology

#### Effective Number of Light $\nu$ Types

“Light” means here with a mass  $<$  about 1 MeV. The quoted values correspond to  $N_{\text{eff}}$ , where  $N_{\text{eff}} = 3.045$  in the Standard Model with  $N_\nu = 3$ . See also reviews on “Big-Bang Nucleosynthesis” and “Neutrinos in Cosmology.”

<u>VALUE</u>	<u>CL%</u>	<u>DOCUMENT ID</u>	<u>TECN</u>	<u>COMMENT</u>
● ● ● We do not use the following data for averages, fits, limits, etc. ● ● ●				
3.12±0.38	95	1 BRIEDEN	22 COSM	BOSS, eBOSS, CMB
2.90±0.15	68	2 KUMAR	22 COSM	BOSS + CMB
2.89±0.14	68	3 YEH	22 COSM	BBN + CMB
2.99±0.17	68	4 AGHANIM	20 COSM	
2.84±0.15	68	5 FIELDS	20 COSM	BBN
2.88±0.17	68	6 IVANOV	20 COSM	Planck and BOSS
2.3–3.2	95	7 VERDE	17 COSM	
2.88±0.16	68	8 CYBURT	16 COSM	BBN
2.88±0.20	95	9 ROSSI	15 COSM	
3.3 ±0.5	95	10 ADE	14 COSM	Planck
3.78 <sup>+0.31</sup> <sub>−0.30</sub>		11 COSTANZI	14 COSM	
3.29±0.31		12 HOU	14 COSM	
$< 3.80$	95	13 LEISTEDT	14 COSM	
$< 4.10$	95	14 MORESCO	12 COSM	
$< 5.79$	95	15 XIA	12 COSM	
$< 4.08$	95	MANGANO	11 COSM	BBN
0.9–8.2		16 ICHIKAWA	07 COSM	
3–7	95	17 CIRELLI	06 COSM	
2.7–4.6	95	18 HANNESTAD	06 COSM	
3.6–7.4	95	17 SELJAK	06 COSM	
$< 4.4$		19 CYBURT	05 COSM	
$< 3.3$		20 BARGER	03C COSM	
1.4–6.8		21 CROTTY	03 COSM	

1.9–6.6	<sup>21</sup> PIERPAOLI	03	COSM
2–4	LISI	99	COSM BBN
< 4.3	OLIVE	99	COSM BBN
< 4.9	COPI	97	Cosmology
< 3.6	HATA	97B	High D/H quasar abs.
< 4.0	OLIVE	97	BBN; high <sup>4</sup> He and <sup>7</sup> Li
< 4.7	CARDALL	96B	COSM High D/H quasar abs.
< 3.9	FIELDS	96	COSM BBN; high <sup>4</sup> He and <sup>7</sup> Li
< 4.5	KERNAN	96	COSM High D/H quasar abs.
< 3.6	OLIVE	95	BBN; $\geq 3$ massless $\nu$
< 3.3	WALKER	91	Cosmology
< 3.4	OLIVE	90	Cosmology
< 4	YANG	84	Cosmology
< 4	YANG	79	Cosmology
< 7	STEIGMAN	77	Cosmology
	PEEBLES	71	Cosmology
<16	<sup>22</sup> SHVARTSMAN	69	Cosmology
	HOYLE	64	Cosmology

<sup>1</sup> BRIEDEN <sup>22</sup> combines large scale structure data from BOSS and eBOSS including the shape of the matter power spectrum with Planck CMB data.

<sup>2</sup> KUMAR <sup>22</sup> combine the reconstructed galaxy power spectrum from BOSS data with Planck CMB data.

<sup>3</sup> YEY <sup>22</sup> combines Planck 2018 CMB data with BBN and observations of deuterium and Helium-4. Supersedes FIELDS <sup>20</sup>.

<sup>4</sup> AGHANIM <sup>20</sup> best fit on number of neutrino types is based on Planck data combined with lensing and baryon acoustic oscillations (BAO). Without BAO, they find  $2.89^{+0.18}_{-0.19}$ . Several other values are quoted using different combinations of data.

<sup>5</sup> FIELDS <sup>20</sup> combines Planck 2018 CMB data with BBN and observations of deuterium and Helium-4.

<sup>6</sup> IVANOV <sup>20</sup> combines 2018 Planck CMB data with baryon acoustic oscillation data from BOSS. This study is based on a full-shape likelihood for the redshift-space galaxy power spectrum of the BOSS data.

<sup>7</sup> Uses Planck Data combined with an independent standard measure of distance to the sound horizon to set a limit on the total number of neutrinos. Only CMB and early-time information are used.

<sup>8</sup> CYBURT <sup>16</sup> combines Planck 2015 CMB data with BBN and observations of deuterium and Helium-4.

<sup>9</sup> ROSSI <sup>15</sup> sets limits on the number of neutrino types using BOSS Lyman alpha forest data combined with Planck CMB data and baryon acoustic oscillations.

<sup>10</sup> Fit to the number of neutrino degrees of freedom from Planck CMB data along with WMAP polarization, high L, and BAO data.

<sup>11</sup> Fit to the number of neutrinos degrees of freedom from Planck CMB data along with BAO, shear and cluster data.

<sup>12</sup> Fit based on the SPT-SZ survey combined with CMB, BAO, and  $H_0$  data.

<sup>13</sup> Constrains the number of neutrino degrees of freedom (marginalizing over the total mass) from CMB, CMB lensing, BAO, and galaxy clustering data.

<sup>14</sup> Limit on the number of light neutrino types from observational Hubble parameter data with seven-year WMAP data, SPT, and the most recent estimate of  $H_0$ . Best fit is  $3.45 \pm 0.65$ .

<sup>15</sup> Limit on the number of light neutrino types from the CFHTLS combined with seven-year WMAP data and a prior on the Hubble parameter. Best fit is  $4.17^{+1.62}_{-1.26}$ . Limit is relaxed to  $3.98^{+2.02}_{-1.20}$  when small scales affected by non-linearities are removed.

- 16 Constrains the number of neutrino types from recent CMB and large scale structure data. No priors on other cosmological parameters are used.
- 17 Constrains the number of neutrino types from recent CMB, large scale structure, Lyman-alpha forest, and SN1a data. The slight preference for  $N_\nu > 3$  comes mostly from the Lyman-alpha forest data.
- 18 Constrains the number of neutrino types from recent CMB and large scale structure data. See also HAMANN 07.
- 19 Limit on the number of neutrino types based on  $^4\text{He}$  and D/H abundance assuming a baryon density fixed to the WMAP data. Limit relaxes to 4.6 if D/H is not used or to 5.8 if only D/H and the CMB are used. See also CYBURT 01 and CYBURT 03.
- 20 Limit on the number of neutrino types based on combination of WMAP data and big-bang nucleosynthesis. The limit from WMAP data alone is 8.3. See also KNELLER 01.  $N_\nu \geq 3$  is assumed to compute the limit.
- 21 95% confidence level range on the number of neutrino flavors from WMAP data combined with other CMB measurements, the 2dfGRS data, and HST data.
- 22 SHVARTSMAN 69 limit inferred from his equations.

### Number Coupling with Less Than Full Weak Strength

VALUE	DOCUMENT ID	TECN
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• • • We do not use the following data for averages, fits, limits, etc. • • •

$<20$	<sup>1</sup> OLIVE	81C COSM
$<20$	<sup>1</sup> STEIGMAN	79 COSM

<sup>1</sup>Limit varies with strength of coupling. See also WALKER 91.

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### REFERENCES FOR Limits on Number of Neutrino Types

BRIEDEN	22	JCAP 2208 024	S. Brieden, Hector Gil-Marín, Licia Verde
KUMAR	22	JCAP 2209 060	S. Kumar, R. Nunes, P.Yadav
YEH	22	JCAP 2210 046	T.-H. Yeh <i>et al.</i> (ILL, MINN)
AGHANIM	20	AA 641 A6	N. Aghanim <i>et al.</i> (Planck Collab.)
FIELDS	20	JCAP 2003 010	B. Fields <i>et al.</i> (ILL, MINN)
Also		JCAP 2011 E02 (errat.)	B. Fields <i>et al.</i> (ILL, MINN)
IVANOV	20	PR D101 083504	M.M. Ivanov, M. Simonic, M. Zaldarriaga (NYU,+)
JANOT	20	PL B803 135319	P. Janot, S. Jadach (CERN, CRAC)
VOUSINAS	20	PL B800 135068	G. Voutsinas <i>et al.</i> (CERN, BOHR)
VERDE	17	JCAP 1704 023	L. Verde <i>et al.</i>
CYBURT	16	RMP 88 015004	R.H. Cyburt <i>et al.</i> (MSU, ILL, MINN)
ROSSI	15	PR D92 063505	G. Rossi <i>et al.</i>
ADE	14	AA 571 A16	P.A.R. Ade <i>et al.</i> (Planck Collab.)
COSTANZI	14	JCAP 1410 081	M. Costanzi <i>et al.</i> (TRST, TRSTI)
HOU	14	APJ 782 74	Z. Hou <i>et al.</i>
LEISTEDT	14	PRL 113 041301	B. Leistedt, H.V. Peiris, L. Verde
MORESCO	12	JCAP 1207 053	M. Moresco <i>et al.</i>
XIA	12	JCAP 1206 010	J.-Q. Xia <i>et al.</i>
MANGANO	11	PL B701 296	G. Mangano, P. Serpico
HAMANN	07	JCAP 0708 021	J. Hamann <i>et al.</i>
ICHIKAWA	07	JCAP 0705 007	K. Ichikawa, M. Kawasaki, F. Takahashi
CIRELLI	06	JCAP 0612 013	M. Cirelli <i>et al.</i>
HANNESTAD	06	JCAP 0611 016	S. Hannestad, G. Raffelt
LEP-SLC	06	PRPL 427 257	ALEPH, DELPHI, L3, OPAL, SLD and working groups
SELJAK	06	JCAP 0610 014	U. Seljak, A. Slosar, P. McDonald
ABDALLAH	05B	EPJ C38 395	J. Abdallah <i>et al.</i> (DELPHI Collab.)
CYBURT	05	ASP 23 313	R.H. Cyburt <i>et al.</i>
ACHARD	04E	PL B587 16	P. Achard <i>et al.</i> (L3 Collab.)
BARGER	03C	PL B566 8	V. Barger <i>et al.</i>
CROTTY	03	PR D67 123005	P. Crotty, J. Lesgourgues, S. Pastor
CYBURT	03	PL B567 227	R.H. Cyburt, B.D. Fields, K.A. Olive
HEISTER	03C	EPJ C28 1	A. Heister <i>et al.</i> (ALEPH Collab.)
PIERPAOLI	03	MNRAS 342 L63	E. Pierpaoli
CYBURT	01	ASP 17 87	R.H. Cyburt, B.D. Fields, K.A. Olive
KNELLER	01	PR D64 123506	J.P. Kneller <i>et al.</i>

ABBIENDI,G	00D	EPJ C18 253	G. Abbiendi <i>et al.</i>	(OPAL Collab.)
ABREU	00Z	EPJ C17 53	P. Abreu <i>et al.</i>	(DELPHI Collab.)
ACCIARRI	99R	PL B470 268	M. Acciarri <i>et al.</i>	(L3 Collab.)
LISI	99	PR D59 123520	E. Lisi, S. Sarkar, F.L. Villante	
OLIVE	99	ASP 11 403	K.A. Olive, D. Thomas	
ABREU	97J	ZPHY C74 577	P. Abreu <i>et al.</i>	(DELPHI Collab.)
COPI	97	PR D55 3389	C.J. Copi, D.N. Schramm, M.S. Turner	(CHIC)
HATA	97B	PR D55 540	N. Hata <i>et al.</i>	(OSU, PENN)
OLIVE	97	ASP 7 27	K.A. Olive, D. Thomas	(MINN, FLOR)
ADAM	96C	PL B380 471	W. Adam <i>et al.</i>	(DELPHI Collab.)
CARDALL	96B	APJ 472 435	C.Y. Cardall, G.M. Fuller	(UCSD)
FIELDS	96	New Ast 1 77	B.D. Fields <i>et al.</i>	(NDAM, CERN, MINN+)
KERNAN	96	PR D54 3681	P.S. Kernan, S. Sarkar	(CASE, OXFTP)
AKERS	95C	ZPHY C65 47	R. Akers <i>et al.</i>	(OPAL Collab.)
OLIVE	95	PL B354 357	K.A. Olive, G. Steigman	(MINN, OSU)
BUSKULIC	93L	PL B313 520	D. Buskulic <i>et al.</i>	(ALEPH Collab.)
LEP	92	PL B276 247	LEP Collabs.	(LEP, ALEPH, DELPHI, L3, OPAL)
WALKER	91	APJ 376 51	T.P. Walker <i>et al.</i>	(HSCA, OSU, CHIC+)
OLIVE	90	PL B236 454	K.A. Olive <i>et al.</i>	(MINN, CHIC, OSU+)
YANG	84	APJ 281 493	J. Yang <i>et al.</i>	(CHIC, BART)
OLIVE	81C	NP B180 497	K.A. Olive, D.N. Schramm, G. Steigman	(CHIC+)
STEIGMAN	79	PRL 43 239	G. Steigman, K.A. Olive, D.N. Schramm	(BART+)
YANG	79	APJ 227 697	J. Yang <i>et al.</i>	(CHIC, YALE, UVA)
STEIGMAN	77	PL 66B 202	G. Steigman, D.N. Schramm, J.E. Gunn	(YALE, CHIC+)
PEEBLES	71	Physical Cosmology	P.Z. Peebles	(PRIN)
		Princeton Univ. Press (1971)		
SHVARTSMAN	69	JETPL 9 184	V.F. Shvartsman	(MOSU)
		Translated from ZETFP 9	315.	
HOYLE	64	NAT 203 1108	F. Hoyle, R.J. Tayler	(CAMB)