Neutrino Mixing

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(A) Neutrino fluxes and event ratios

Events (observed/expected) from accelerator ν_μ experiments.

Some neutrino oscillation experiments compare the flux in two or more detectors. This is usually quoted as the ratio of the event rate in the far detector to the expected rate based on an extrapolation from the near detector in the absence of oscillations.

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
0.71 ± 0.08	[1] AHN	06A	K2K	K2K to Super-K
0.64 ± 0.05	[2] MICHAEL	06	MINS	All charged current events
$0.71^{+0.08}_{-0.09}$	[3] ALIU	05	K2K	KEK to Super-K
$0.70^{+0.10}_{-0.11}$	[4] AHN	03	K2K	KEK to Super-K

[1] Based on the observation of 112 events when $158.1^{+9.2}_{-8.6}$ were expected without oscillations. Including not only the number of events but also the shape of the energy distribution, the evidence for oscillation is at the level of about 4.3 σ. Supersedes ALIU 05.

[2] This ratio is based on the observation of 215 events compared to an expectation of 336 ± 14 without oscillations.

[3] This ratio is based on the observation of 107 events at the far detector 250 km away from KEK, and an expectation of 151^{+12}_{-10}.

[4] This ratio is based on the observation of 56 events with an expectation of $80.1^{+6.2}_{-5.4}$.

Events (observed/expected) from reactor $\overline{\nu}_e$ experiments.

The quoted values are the ratios of the measured reactor $\overline{\nu}_e$ event rate at the quoted distances, and the rate expected without oscillations. The expected rate is based on the experimental data for the most significant reactor fuels (^{235}U, ^{239}Pu, ^{241}Pu) and on calculations for ^{238}U.

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$0.658\pm0.044\pm0.047$	[5] ARAKI	05	KLND	Japanese react. ~ 180 km
$0.611\pm0.085\pm0.041$	[6] EGUCHI	03	KLND	Japanese react. ~ 180 km
$1.01\ \pm0.024\pm0.053$	[7] BOEHM	01		Palo Verde react. 0.75–0.89 km
$1.01\ \pm0.028\pm0.027$	[8] APOLLONIO	99	CHOZ	Chooz reactors 1 km
$0.987\pm0.006\pm0.037$	[9] GREENWOOD	96		Savannah River, 18.2 m
$0.988\pm0.004\pm0.05$	ACHKAR	95	CNTR	Bugey reactor, 15 m
$0.994\pm0.010\pm0.05$	ACHKAR	95	CNTR	Bugey reactor, 40 m
$0.915\pm0.132\pm0.05$	ACHKAR	95	CNTR	Bugey reactor, 95 m
$0.987\pm0.014\pm0.027$	[10] DECLAIS	94	CNTR	Bugey reactor, 15 m
$0.985\pm0.018\pm0.034$	KUVSHINN...	91	CNTR	Rovno reactor
$1.05\ \pm0.02\ \pm0.05$	VUILLEUMIER	82		Gösgen reactor
$0.955\pm0.035\pm0.110$	[11] KWON	81		$\overline{\nu}_e p \to e^+ n$
$0.89\ \pm0.15$	[11] BOEHM	80		$\overline{\nu}_e p \to e^+ n$

[5] Updated result of KamLAND, including the data used in EGUCHI 03. Note that the survival probabilities for different periods are not directly comparable because the effective baseline varies with power output of the reactor sources involved, and there were large variations in the reactor power production in Japan in 2003.

[6] EGUCHI 03 observe reactor neutrino disappearance at $\sim$ 180 km baseline to various Japanese nuclear power reactors.

[7] BOEHM 01 search for neutrino oscillations at 0.75 and 0.89 km distance from the Palo Verde reactors.

[8] APOLLONIO 99, APOLLONIO 98 search for neutrino oscillations at 1.1 km fixed distance from Chooz reactors. They use $\overline{\nu}_e p \rightarrow e^+ n$ in Gd-loaded scintillator target. APOLLONIO 99 supersedes APOLLONIO 98. See also APOLLONIO 03 for detailed description.

[9] GREENWOOD 96 search for neutrino oscillations at 18 m and 24 m from the reactor at Savannah River.

[10] DECLAIS 94 result based on integral measurement of neutrons only. Result is ratio of measured cross section to that expected in standard *V-A* theory. Replaced by ACHKAR 95.

[11] KWON 81 represents an analysis of a larger set of data from the same experiment as BOEHM 80.

Atmospheric neutrinos

Neutrinos and antineutrinos produced in the atmosphere induce μ-like and e-like events in underground detectors. The ratio of the numbers of the two kinds of events is defined as μ/e. It has the advantage that systematic effects, such as flux uncertainty, tend to cancel, for both experimental and theoretical values of the ratio. The "ratio of the ratios" of experimental to theoretical μ/e, $R(\mu/e)$, or that of experimental to theoretical μ/total, $R(\mu/\text{total})$ with total = μ+e, is reported below. If the actual value is not unity, the value obtained in a given experiment may depend on the experimental conditions. In addition, the measured "up-down asymmetry" for μ ($N_{up}(\mu)/N_{down}(\mu)$) or e ($N_{up}(e)/N_{down}(e)$) is reported. The expected "up-down asymmetry" is nearly unity if there is no neutrino oscillation.

R(μ/e) = (Measured Ratio μ/e) / (Expected Ratio μ/e)

VALUE	*DOCUMENT ID*	*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •			
$0.658\pm0.016\pm0.035$	[12] ASHIE 05	SKAM	sub-GeV
$0.702^{+0.032}_{-0.030}\pm0.101$	[13] ASHIE 05	SKAM	multi-GeV
$0.69\pm0.10\pm0.06$	[14] SANCHEZ 03	SOU2	Calorimeter raw data
	[15] FUKUDA 96B	KAMI	Water Cherenkov
$1.00\pm0.15\pm0.08$	[16] DAUM 95	FREJ	Calorimeter
$0.60^{+0.06}_{-0.05}\pm0.05$	[17] FUKUDA 94	KAMI	sub-GeV
$0.57^{+0.08}_{-0.07}\pm0.07$	[18] FUKUDA 94	KAMI	multi-Gev
	[19] BECKER-SZ... 92B	IMB	Water Cherenkov

[12] ASHIE 05 results are based on an exposure of 92 kton yr during the complete Super-Kamiokande I running period. The analyzed data sample consists of fully-contained single-ring e-like events with 0.1 GeV/c $< p_e$ and μ-like events 0.2 GeV/c $< p_\mu$, both having a visible energy $<$ 1.33 GeV. These criteria match the definition used by FUKUDA 94.

[13] ASHIE 05 results are based on an exposure of 92 kton yr during the complete Super-Kamiokande I running period. The analyzed data sample consists of fully-contained single-ring events with visible energy $>$ 1.33 GeV and partially-contained events. All partially-contained events are classified as μ-like.

[14] SANCHEZ 03 result is based on an exposure of 5.9 kton yr, and updates ALLISON 99 result. The analyzed data sample consists of fully-contained e-flavor and μ-flavor events having lepton momentum $>$ 0.3 GeV/c.

[15] FUKUDA 96B studied neutron background in the atmospheric neutrino sample observed in the Kamiokande detector. No evidence for the background contamination was found.

[16] DAUM 95 results are based on an exposure of 2.0 kton yr which includes the data used by BERGER 90B. This ratio is for the contained and semicontained events. DAUM 95 also report $R(\mu/e) = 0.99 \pm 0.13 \pm 0.08$ for the total neutrino induced data sample which includes upward going stopping muons and horizontal muons in addition to the contained and semicontained events.

[17] FUKUDA 94 result is based on an exposure of 7.7 kton yr and updates the HIRATA 92 result. The analyzed data sample consists of fully-contained e-like events with $0.1 < p_e < 1.33$ GeV/c and fully-contained μ-like events with $0.2 < p_\mu < 1.5$ GeV/c.

[18] FUKUDA 94 analyzed the data sample consisting of fully contained events with visible energy $>$ 1.33 GeV and partially contained μ-like events.

[19] BECKER-SZENDY 92B reports the fraction of nonshowering events (mostly muons from atmospheric neutrinos) as $0.36 \pm 0.02 \pm 0.02$, as compared with expected fraction $0.51 \pm 0.01 \pm 0.05$. After cutting the energy range to the Kamiokande limits, BEIER 92 finds $R(\mu/e)$ very close to the Kamiokande value.

R(ν_μ) = (Measured Flux of ν_μ) / (Expected Flux of ν_μ)

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
0.84 ± 0.12	[20] ADAMSON	06	MINS	MINOS atmospheric
$0.72 \pm 0.026 \pm 0.13$	[21] AMBROSIO	01	MCRO	upward through-going
$0.57 \pm 0.05 \pm 0.15$	[22] AMBROSIO	00	MCRO	upgoing partially contained
$0.71 \pm 0.05 \pm 0.19$	[23] AMBROSIO	00	MCRO	downgoing partially contained + upgoing stopping
$0.74 \pm 0.036 \pm 0.046$	[24] AMBROSIO	98	MCRO	Streamer tubes
	[25] CASPER	91	IMB	Water Cherenkov
	[26] AGLIETTA	89	NUSX	
0.95 ± 0.22	[27] BOLIEV	81		Baksan
0.62 ± 0.17	CROUCH	78		Case Western/UCI

[20] ADAMSON 06 uses a measurement of 107 total neutrinos compared to an expected rate of 127 ± 13 without oscillations.

[21] AMBROSIO 01 result is based on the upward through-going muon tracks with $E_\mu > 1$ GeV. The data came from three different detector configurations, but the statistics is largely dominated by the full detector run, from May 1994 to December 2000. The total live time, normalized to the full detector configuration, is 6.17 years. The first error is the statistical error, the second is the systematic error, dominated by the theoretical error in the predicted flux.

[22] AMBROSIO 00 result is based on the upgoing partially contained event sample. It came from 4.1 live years of data taking with the full detector, from April 1994 to February 1999. The average energy of atmospheric muon neutrinos corresponding to this sample is 4 GeV. The first error is statistical, the second is the systematic error, dominated by

the 25% theoretical error in the rate (20% in the flux and 15% in the cross section, added in quadrature). Within statistics, the observed deficit is uniform over the zenith angle.

[23] AMBROSIO 00 result is based on the combined samples of downgoing partially contained events and upgoing stopping events. These two subsamples could not be distinguished due to the lack of timing information. The result came from 4.1 live years of data taking with the full detector, from April 1994 to February 1999. The average energy of atmospheric muon neutrinos corresponding to this sample is 4 GeV. The first error is statistical, the second is the systematic error, dominated by the 25% theoretical error in the rate (20% in the flux and 15% in the cross section, added in quadrature). Within statistics, the observed deficit is uniform over the zenith angle.

[24] AMBROSIO 98 result is for all nadir angles and updates AHLEN 95 result. The lower cutoff on the muon energy is 1 GeV. In addition to the statistical and systematic errors, there is a Monte Carlo flux error (theoretical error) of ±0.13. With a neutrino oscillation hypothesis, the fit either to the flux or zenith distribution independently yields $\sin^2 2\theta = 1.0$ and $\Delta(m^2) \sim$ a few times 10^{-3} eV^2. However, the fit to the observed zenith distribution gives a maximum probability for χ^2 of only 5% for the best oscillation hypothesis.

[25] CASPER 91 correlates showering/nonshowering signature of single-ring events with parent atmospheric-neutrino flavor. They find nonshowering ($\approx \nu_\mu$ induced) fraction is $0.41 \pm 0.03 \pm 0.02$, as compared with expected 0.51 ± 0.05 (syst).

[26] AGLIETTA 89 finds no evidence for any anomaly in the neutrino flux. They define ρ = (measured number of ν_e's)/(measured number of ν_μ's). They report ρ(measured)=ρ(expected) = $0.96^{+0.32}_{-0.28}$.

[27] From this data BOLIEV 81 obtain the limit $\Delta(m^2) \leq 6 \times 10^{-3}$ eV^2 for maximal mixing, $\nu_\mu \not\rightarrow \nu_\mu$ type oscillation.

R(μ/total) = (Measured Ratio μ/total) / (Expected Ratio μ/total)

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$1.1^{+0.07}_{-0.12} \pm 0.11$	[28] CLARK	97	IMB	multi-GeV

[28] CLARK 97 obtained this result by an analysis of fully contained and partially contained events in the IMB water-Cherenkov detector with visible energy > 0.95 GeV.

$N_{\mathrm{up}}(\mu)/N_{\mathrm{down}}(\mu)$

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$0.551^{+0.035}_{-0.033} \pm 0.004$	[29] ASHIE	05	SKAM	multi-GeV

[29] ASHIE 05 results are based on an exposure of 92 kton yr during the complete Super-Kamiokande I running period. The analyzed data sample consists of fully-contained single-ring μ-like events with visible energy > 1.33 GeV and partially-contained events. All partially-contained events are classified as μ-like. Upward-going events are those with $-1 < \cos(\text{zenith angle}) < -0.2$ and downward-going events are those with $0.2 < \cos(\text{zenith angle}) < 1$. The μ-like up-down ratio for the multi-GeV data deviates from 1 (the expectation for no atmospheric ν_μ oscillations) by more than 12 standard deviations.

$N_{\mathrm{up}}(e)/N_{\mathrm{down}}(e)$

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$0.961^{+0.086}_{-0.079} \pm 0.016$	[30] ASHIE	05	SKAM	multi-GeV

[30] ASHIE 05 results are based on an exposure of 92 kton yr during the complete Super-Kamiokande I running period. The analyzed data sample consists of fully-contained single-ring *e*-like events with visible energy > 1.33 GeV. Upward-going events are those with $-1 < \cos(\text{zenith angle}) < -0.2$ and downward-going events are those with $0.2 < \cos(\text{zenith angle}) < 1$. The *e*-like up-down ratio for the multi-GeV data is consistent with 1 (the expectation for no atmospheric ν_e oscillations).

R(up/down; μ) = (Measured up/down; μ) / (Expected up/down; μ)

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$0.62^{+0.19}_{-0.14} \pm 0.02$	[31] ADAMSON	06	MINS	atmospheric ν with far detector

[31] ADAMSON 06 result is obtained with the MINOS far detector with an exposure of 4.54 kton yr. The expected ratio is calculated with no neutrino oscillation.

R(μ^+/μ^-) = (Measured N(μ^+)/N(μ^-)) / (Expected N(μ^+)/N(μ^-))

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$0.96^{+0.38}_{-0.27} \pm 0.15$	[32] ADAMSON	06	MINS	atmospheric ν with far detector

[32] ADAMSON 06 result is obtained with the MINOS far detector with an exposure of 4.54 kton yr. The expected ratio is calculated by assuming the same oscillation parameters for neutrinos and antineutrinos.

Solar neutrinos

Solar neutrinos are produced by thermonuclear fusion reactions in the Sun. Radiochemical experiments measure particular combinations of fluxes from various neutrino-producing reactions, whereas water-Cherenkov experiments mainly measure a flux of neutrinos from decay of ^{8}B. Solar neutrino fluxes are composed of all active neutrino species, ν_e, ν_μ, and ν_τ. In addition, some other mechanisms may cause antineutrino components in solar neutrino fluxes. Each measurement method is sensitive to a particular component or a combination of components of solar neutrino fluxes. For details, see the following minireview.

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ν_e Capture Rates from Radiochemical Experiments

1 SNU (Solar Neutrino Unit) = 10^{-36} captures per atom per second.

VALUE (SNU)	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$62.9 \; ^{+5.5}_{-5.3} \; \pm 2.5$	[33] ALTMANN	05	GNO	$^{71}\text{Ga} \rightarrow {}^{71}\text{Ge}$
$69.3 \; \pm 4.1 \; \pm 3.6$	[34] ALTMANN	05	GNO	GNO + GALX combined
$70.8 \; ^{+5.3}_{-5.2} \; ^{+3.7}_{-3.2}$	[35] ABDURASHI...	02	SAGE	$^{71}\text{Ga} \rightarrow {}^{71}\text{Ge}$
$77.5 \; \pm 6.2 \; ^{+4.3}_{-4.7}$	[36] HAMPEL	99	GALX	$^{71}\text{Ga} \rightarrow {}^{71}\text{Ge}$
$2.56 \pm 0.16 \pm 0.16$	[37] CLEVELAND	98	HOME	$^{37}\text{Cl} \rightarrow {}^{37}\text{Ar}$

[33] ALTMANN 05 reports the complete result from the GNO solar neutrino experiment (GNO I+II+III), which is the successor project of GALLEX. Experimental technique of GNO is essentially the same as that of GALLEX. The run data cover the period 20 May 1998 through 9 April 2003.

[34] Combined result of GALLEX I+II+III+IV (HAMPEL 99) and GNO I+II+III.

[35] ABDURASHITOV 02 report a combined analysis of 92 runs of the SAGE solar-neutrino experiment during the period January 1990 through December 2001, and updates the ABDURASHITOV 99B result. A total of 406.4 ^{71}Ge events were observed. No evidence was found for temporal variations of the neutrino capture rate over the entire observation period.

[36] HAMPEL 99 report the combined result for GALLEX I+II+III+IV (65 runs in total), which update the HAMPEL 96 result. The GALLEX IV result (12 runs) is $118.4 \pm 17.8 \pm 6.6$ SNU. (HAMPEL 99 discuss the consistency of partial results with the mean.) The GALLEX experimental program has been completed with these runs. The total run data cover the period 14 May 1991 through 23 January 1997. A total of 300 ^{71}Ge events were observed.

[37] CLEVELAND 98 is a detailed report of the ^{37}Cl experiment at the Homestake Mine. The average solar neutrino-induced ^{37}Ar production rate from 108 runs between 1970 and 1994 updates the DAVIS 89 result.

ϕ_{ES} (^{8}B)

^{8}B solar-neutrino flux measured via νe elastic scattering. This process is sensitive to all active neutrino flavors, but with reduced sensitivity to ν_{μ}, ν_{τ} due to the cross-section difference, $\sigma(\nu_{\mu,\tau} e) \sim 0.16\sigma(\nu_e e)$. If the ^{8}B solar-neutrino flux involves nonelectron flavor active neutrinos, their contribution to the flux is $\sim$ 0.16 times of ν_e.

VALUE (10^6 cm^{-2}s^{-1})	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$2.35 \pm 0.02 \pm 0.08$	[38] HOSAKA	06	SKAM	average flux
$2.35 \pm 0.22 \pm 0.15$	[39] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape not constrained
$2.34 \pm 0.23 ^{+0.15}_{-0.14}$	[39] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape constrained
$2.39 ^{+0.24}_{-0.23} \pm 0.12$	[40] AHMAD	02	SNO	average flux
$2.39 \pm 0.34 ^{+0.16}_{-0.14}$	[41] AHMAD	01	SNO	average flux
$2.80 \pm 0.19 \pm 0.33$	[42] FUKUDA	96	KAMI	average flux
2.70 ± 0.27	[42] FUKUDA	96	KAMI	day flux
$2.87 ^{+0.27}_{-0.26}$	[42] FUKUDA	96	KAMI	night flux

[38] HOSAKA 06 reports the final results for 1496 live days with Super-Kamiokande-I between May 31, 1996 and July 15, 2001, and replace FUKUDA 02 results. The analysis threshold is 5 MeV except for the first 280 live days (6.5 MeV).

[39] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, corresponding to 391.4 live days, and update AHMED 04A. The *CC*, *ES*, and *NC* events were statistically separated. In one method, the ^{8}B energy spectrum was not constrained. In the other method, the constraint of an undistorted ^{8}B energy spectrum was added for comparison with AHMAD 02 results.

[40] AHMAD 02 reports the ^{8}B solar-neutrino flux measured via νe elastic scattering above the kinetic energy threshold of 5 MeV. The data correspond to 306.4 live days with SNO between November 2, 1999 and May 28, 2001, and updates AHMAD 01 results.

[41] AHMAD 01 reports the ^{8}B solar-neutrino flux measured via νe elastic scattering above the kinetic energy threshold of 6.75 MeV. The data correspond to 241 live days with SNO between November 2, 1999 and January 15, 2001.

[42] FUKUDA 96 results are for a total of 2079 live days with Kamiokande II and III from January 1987 through February 1995, covering the entire solar cycle 22, with threshold $E_e > 9.3$ MeV (first 449 days), > 7.5 MeV (middle 794 days), and > 7.0 MeV (last 836 days). These results update the HIRATA 90 result for the average ^{8}B solar-neutrino flux and HIRATA 91 result for the day-night variation in the ^{8}B solar-neutrino flux. The total data sample was also analyzed for short-term variations: within experimental errors, no strong correlation of the solar-neutrino flux with the sunspot numbers was found.

ϕ_{CC} (^{8}B)

^{8}B solar-neutrino flux measured with charged-current reaction which is sensitive exclusively to ν_e.

VALUE (10^6 cm^{-2}s^{-1})	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$1.68 \pm 0.06 {}^{+0.08}_{-0.09}$	[43] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape not const.
$1.72 \pm 0.05 \pm 0.11$	[43] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape constrained
$1.76 {}^{+0.06}_{-0.05} \pm 0.09$	[44] AHMAD	02	SNO	average flux
$1.75 \pm 0.07 {}^{+0.12}_{-0.11} \pm 0.05$	[45] AHMAD	01	SNO	average flux

[43] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, corresponding to 391.4 live days, and update AHMED 04A. The *CC*, *ES*, and *NC* events were statistically separated. In one method, the ^{8}B energy spectrum was not constrained. In the other method, the constraint of an undistorted ^{8}B energy spectrum was added for comparison with AHMAD 02 results.

[44] AHMAD 02 reports the SNO result of the ^{8}B solar-neutrino flux measured with charged-current reaction on deuterium, $\nu_e d \rightarrow ppe^-$, above the kinetic energy threshold of 5 MeV. The data correspond to 306.4 live days with SNO between November 2, 1999 and May 28, 2001, and updates AHMAD 01 results.

[45] AHMAD 01 reports the first SNO result of the ^{8}B solar-neutrino flux measured with the charged-current reaction on deuterium, $\nu_e d \rightarrow ppe^-$, above the kinetic energy threshold of 6.75 MeV. The data correspond to 241 live days with SNO between November 2, 1999 and January 15, 2001.

ϕ_{NC} (^{8}B)

^{8}B solar neutrino flux measured with neutral-current reaction, which is equally sensitive to ν_e, ν_μ, and ν_τ.

VALUE (10^6 cm^{-2}s^{-1})	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$4.94 \pm 0.21 {}^{+0.38}_{-0.34}$	[46] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape not const.
$4.81 \pm 0.19 {}^{+0.28}_{-0.27}$	[46] AHARMIM	05A	SNO	Salty D_2O; ^{8}B shape constrained
$5.09 {}^{+0.44}_{-0.43} {}^{+0.46}_{-0.43}$	[47] AHMAD	02	SNO	average flux; ^{8}B shape const.
$6.42 \pm 1.57 {}^{+0.55}_{-0.58}$	[47] AHMAD	02	SNO	average flux; ^{8}B shape not const.

[46] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, corresponding to 391.4 live days, and update AHMED 04A. The *CC*, *ES*, and *NC* events were statistically separated. In one method, the ^{8}B energy spectrum was not constrained. In the other method, the constraint of an undistorted ^{8}B energy spectrum was added for comparison with AHMAD 02 results.

[47] AHMAD 02 reports the first SNO result of the ^{8}B solar-neutrino flux measured with the neutral-current reaction on deuterium, $\nu_\ell d \rightarrow n p \nu_\ell$, above the neutral-current reaction threshold of 2.2 MeV. The data correspond to 306.4 live days with SNO between November 2, 1999 and May 28, 2001.

$\phi_{\nu_\mu + \nu_\tau}$ (^{8}B)

Nonelectron-flavor active neutrino component (ν_μ and ν_τ) in the ^{8}B solar-neutrino flux.

VALUE (10^6 cm^{-2}s^{-1})	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$3.26 \pm 0.25 {}^{+0.40}_{-0.35}$	[48] AHARMIM	05A	SNO	From ϕ_{NC}, ϕ_{CC}, and ϕ_{ES}; ^{8}B shape not const.
$3.09 \pm 0.22 {}^{+0.30}_{-0.27}$	[48] AHARMIM	05A	SNO	From ϕ_{NC}, ϕ_{CC}, and ϕ_{ES}; ^{8}B shape constrained
$3.41 \pm 0.45 {}^{+0.48}_{-0.45}$	[49] AHMAD	02	SNO	From ϕ_{NC}, ϕ_{CC}, and ϕ_{ES}
3.69 ± 1.13	[50] AHMAD	01		Derived from SNO+SuperKam, water Cherenkov

[48] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, corresponding to 391.4 live days, and update AHMED 04A. The *CC*, *ES*, and *NC* events were statistically separated. In one method, the ^{8}B energy spectrum was not constrained. In the other method, the constraint of an undistorted ^{8}B energy spectrum was added for comparison with AHMAD 02 results.

[49] AHMAD 02 deduced the nonelectron-flavor active neutrino component (ν_μ and ν_τ) in the ^{8}B solar-neutrino flux, by combining the charged-current result, the νe elastic-scattering result and the neutral-current result.

[50] AHMAD 01 deduced the nonelectron-flavor active neutrino component (ν_μ and ν_τ) in the ^{8}B solar-neutrino flux, by combining the SNO charged-current result (AHMAD 01) and the Super-Kamiokande νe elastic-scattering result (FUKUDA 01).

Total Flux of Active ^{8}B Solar Neutrinos

Total flux of active neutrinos (ν_e, ν_μ, and ν_τ).

VALUE (10^6 cm^{-2}s^{-1})	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
$4.94 \pm 0.21 {}^{+0.38}_{-0.34}$	[51] AHARMIM	05A	SNO	From ϕ_{NC}; ^{8}B shape not const.
$4.81 \pm 0.19 {}^{+0.28}_{-0.27}$	[51] AHARMIM	05A	SNO	From ϕ_{NC}; ^{8}B shape constrained

$5.09^{+0.44}_{-0.43}{}^{+0.46}_{-0.43}$	[52] AHMAD	02	SNO	Direct measurement from ϕ_{NC}
5.44 ± 0.99	[53] AHMAD	01		Derived from SNO+SuperKam, water Cherenkov

[51] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, corresponding to 391.4 live days, and update AHMED 04A. The *CC*, *ES*, and *NC* events were statistically separated. In one method, the ^{8}B energy spectrum was not constrained. In the other method, the constraint of an undistorted ^{8}B energy spectrum was added for comparison with AHMAD 02 results.

[52] AHMAD 02 determined the total flux of active ^{8}B solar neutrinos by directly measuring the neutral-current reaction, $\nu_\ell d \rightarrow np\nu_\ell$, which is equally sensitive to ν_e, ν_μ, and ν_τ.

[53] AHMAD 01 deduced the total flux of active ^{8}B solar neutrinos by combining the SNO charged-current result (AHMAD 01) and the Super-Kamiokande νe elastic-scattering result (FUKUDA 01).

Day-Night Asymmetry (^{8}B)

$A = (\phi_{\text{night}} - \phi_{\text{day}}) / \phi_{\text{average}}$

VALUE	*DOCUMENT ID*	*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •			
$0.021 \pm 0.020^{+0.012}_{-0.013}$	[54] HOSAKA 06	SKAM	Based on ϕ_{ES}
$0.017 \pm 0.016^{+0.012}_{-0.013}$	[55] HOSAKA 06	SKAM	Fitted in the LMA region
$-0.056 \pm 0.074 \pm 0.053$	[56] AHARMIM 05A	SNO	From salty SNO ϕ_{CC}
$-0.037 \pm 0.063 \pm 0.032$	[56] AHARMIM 05A	SNO	From salty SNO ϕ_{CC}; const. of no ϕ_{NC} asymmetry
$0.14 \pm 0.063^{+0.015}_{-0.014}$	[57] AHMAD 02B	SNO	Derived from SNO ϕ_{CC}
$0.07 \pm 0.049^{+0.013}_{-0.012}$	[58] AHMAD 02B	SNO	Const. of no ϕ_{NC} asymmetry

[54] HOSAKA 06 reports the final results for 1496 live days with Super-Kamiokande-I between May 31, 1996 and July 15, 2001, and replace FUKUDA 02 results. The analysis threshold is 5 MeV except for the first 280 live days (6.5 MeV).

[55] This result with reduced statistical uncertainty is obtained by assuming two-neutrino oscillations within the LMA (large mixing angle) region and by fitting the time variation of the solar neutrino flux measured via ν_e elastic scattering to the variations expected from neutrino oscillations. For details, see SMY 04. There is an additional small systematic error of ±0.0004 coming from uncertainty of oscillation parameters.

[56] AHARMIM 05A measurements were made with dissolved NaCl (0.195% by weight) in heavy water over the period between July 26, 2001 and August 28, 2003, with 176.5 days of the live time recorded during the day and 214.9 days during the night. This result is obtained with the spectral distribution of the CC events not constrained to the ^{8}B shape.

[57] AHMAD 02B results are based on the charged-current interactions recorded between November 2, 1999 and May 28, 2001, with the day and night live times of 128.5 and 177.9 days, respectively.

[58] AHMAD 02B results are derived from the charged-current interactions, neutral-current interactions, and νe elastic scattering, with the total flux of active neutrinos constrained to have no asymmetry. The data were recorded between November 2, 1999 and May 28, 2001, with the day and night live times of 128.5 and 177.9 days, respectively.

ϕ_{ES} (hep)

hep solar-neutrino flux measured via νe elastic scattering. This process is sensitive to all active neutrino flavors, but with reduced sensitivity to ν_μ, ν_τ due to the cross-section difference, $\sigma(\nu_{\mu,\tau} e) \sim 0.16\sigma(\nu_e e)$. If the hep solar-neutrino flux involves nonelectron flavor active neutrinos, their contribution to the flux is $\sim$ 0.16 times of ν_e.

VALUE (10^3 cm^{-2}s^{-1})	*CL%*	*DOCUMENT ID*	*TECN*
• • • We do not use the following data for averages, fits, limits, etc. • • •			
<73	90	[59] HOSAKA 06	SKAM

[59] HOSAKA 06 result is obtained from the recoil electron energy window of 18–21 MeV, and updates FUKUDA 01 result.

$\phi_{\overline{\nu}_e}$ (^{8}B)

Searches are made for electron antineutrino flux from the Sun. Flux limits listed here are derived relative to the BS05(OP) Standard Solar Model ^{8}B solar neutrino flux, with an assumption that solar $\overline{\nu}_e$s follow an unoscillated ^{8}B neutrino spectrum.

VALUE (%)	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<1.9	90	[60] BALATA 06	CNTR	$1.8< E_{\overline{\nu}_e} < 20.0$ MeV
<0.72	90	AHARMIM 04	SNO	$4.0< E_{\overline{\nu}_e} < 14.8$ MeV
<0.025	90	EGUCHI 04	KLND	$8.3< E_{\overline{\nu}_e} < 14.8$ MeV
<0.7	90	GANDO 03	SKAM	$8.0< E_{\overline{\nu}_e} < 20.0$ MeV
<1.7	90	AGLIETTA 96	LSD	$7< E_{\overline{\nu}_e} < 17$ MeV

[60] BALATA 06 obtained this result from the search for $\overline{\nu}_e$ interactions with Counting Test Facility (the prototype of the Borexino detector).

(B) Three-neutrino mixing parameters

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$\sin^2(2\theta_{12})$

VALUE	*DOCUMENT ID*	*TECN*	*COMMENT*
$0.86^{+0.03}_{-0.04}$	[61] AHARMIM 05A	FIT	KamLAND + global solar
• • • We do not use the following data for averages, fits, limits, etc. • • •			
$0.85^{+0.04}_{-0.06}$	[62] HOSAKA 06	FIT	KamLAND + global solar
$0.85^{+0.06}_{-0.05}$	[63] HOSAKA 06	FIT	SKAM+SNO+KamLAND
$0.86^{+0.05}_{-0.07}$	[64] HOSAKA 06	FIT	SKAM+SNO
0.75–0.95	[65] AHARMIM 05A	FIT	global solar
0.82±0.05	[66] ARAKI 05	FIT	KamLAND + global solar
0.82±0.04	[67] AHMED 04A	FIT	KamLAND + global solar

0.71–0.93	[68] AHMED	04A	FIT	global solar
$0.85^{+0.05}_{-0.07}$	[69] SMY	04	FIT	KamLAND + global solar
$0.83^{+0.06}_{-0.08}$	[70] SMY	04	FIT	global solar
$0.87^{+0.07}_{-0.08}$	[71] SMY	04	FIT	SKAM + SNO
0.62–0.88	[72] AHMAD	02B	FIT	global solar
0.62–0.95	[73] FUKUDA	02	FIT	global solar

[61] The result given by AHARMIM 05A is $\theta = (33.9 \pm 1.6)^\circ$. This result is obtained by a two-neutrino oscillation analysis using SNO pure deuteron and salt phase data, SK ν_e data, Cl and Ga CC data, and KamLAND data (ARAKI 05). *CPT* invariance is assumed. AHARMIM 05A also quotes $\theta = (33.9^{+2.4}_{-2.2})^\circ$ as the error enveloping the 68% CL two-dimensional region. This translates into $\sin^2 2\,\theta = 0.86^{+0.05}_{-0.06}$.

[62] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using SK ν_e data, CC data from other solar neutrino experiments, and KamLAND data (ARAKI 05). *CPT* invariance is assumed.

[63] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using the data from Super-Kamiokande, SNO (AHMAD 02 and AHMAD 02B), and KamLAND (ARAKI 05) experiments. *CPT* invariance is assumed.

[64] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using the Super-Kamiokande and SNO (AHMAD 02 and AHMAD 02B) solar neutrino data.

[65] AHARMIM 05A obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in figure 35a of AHARMIM 05A. AHARMIM 05A also quotes $\tan^2\theta = 0.45^{+0.09}_{-0.08}$ as the error enveloping the 68% CL two-dimensional region. This translates into $\sin^2 2\,\theta = 0.86^{+0.05}_{-0.07}$.

[66] ARAKI 05 obtained this result by a two-neutrino oscillation analysis using KamLAND and solar neutrino data. *CPT* invariance is assumed. The 1σ error shown here is translated from the number provided by the KamLAND collaboration, $\tan^2\theta = 0.40^{+0.07}_{-0.05}$. The corresponding number quoted in ARAKI 05 is $\tan^2\theta = 0.40^{+0.10}_{-0.07}$ ($\sin^2 2\,\theta = 0.82 \pm 0.07$), which envelops the 68% CL two-dimensional region.

[67] The result given by AHMED 04A is $\theta = (32.5^{+1.7}_{-1.6})^\circ$. This result is obtained by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (EGUCHI 03). *CPT* invariance is assumed. AHMED 04A also quotes $\theta = (32.5^{+2.4}_{-2.3})^\circ$ as the error enveloping the 68% CL two-dimensional region. This translates into $\sin^2 2\,\theta = 0.82 \pm 0.06$.

[68] AHMED 04A obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 5(a) of AHMED 04A. The best-fit point is $\Delta(m^2) = 6.5 \times 10^{-5}$ eV^2, $\tan^2\theta = 0.40$ ($\sin^2 2\,\theta = 0.82$).

[69] The result given by SMY 04 is $\tan^2\theta = 0.44 \pm 0.08$. This result is obtained by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (IANNI 03). *CPT* invariance is assumed.

[70] SMY 04 obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The 1σ errors are read from Fig. 6(a) of SMY 04.

[71] SMY 04 obtained this result by a two-neutrino oscillation analysis using the Super-Kamiokande and SNO (AHMAD 02 and AHMAD 02B) solar neutrino data. The 1σ errors are read from Fig. 6(a) of SMY 04.

[72] AHMAD 02B obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 4(b) of AHMAD 02B. The best fit point is $\Delta(m^2) = 5.0 \times 10^{-5}$ eV^2 and $\tan\theta = 0.34$ ($\sin^2 2\,\theta = 0.76$).

[73] FUKUDA 02 obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 4 of FUKUDA 02. The best fit point is $\Delta(m^2) = 6.9 \times 10^{-5}$ eV^2 and $\tan^2\theta = 0.38$ ($\sin^2 2\,\theta = 0.80$).

Δm^2_{21}

VALUE (10^{-5} eV^2)	*DOCUMENT ID*		*TECN*	*COMMENT*
8.0±0.3	[74] HOSAKA	06	FIT	KamLAND + global solar

• • • We do not use the following data for averages, fits, limits, etc. • • •

VALUE (10^{-5} eV^2)	*DOCUMENT ID*		*TECN*	*COMMENT*
8.0±0.3	[75] HOSAKA	06	FIT	SKAM+SNO+KamLAND
$6.3^{+3.7}_{-1.5}$	[76] HOSAKA	06	FIT	SKAM+SNO
5–12	[77] HOSAKA	06	FIT	SKAM day/night in the LMA region
$8.0^{+0.4}_{-0.3}$	[78] AHARMIM	05A	FIT	KamLAND + global solar LMA
3.3–14.4	[79] AHARMIM	05A	FIT	global solar
$7.9^{+0.4}_{-0.3}$	[80] ARAKI	05	FIT	KamLAND + global solar
$7.1^{+1.0}_{-0.3}$	[81] AHMED	04A	FIT	KamLAND + global solar
3.2–13.7	[82] AHMED	04A	FIT	global solar
$7.1^{+0.6}_{-0.5}$	[83] SMY	04	FIT	KamLAND + global solar
$6.0^{+1.7}_{-1.6}$	[84] SMY	04	FIT	global solar
$6.0^{+2.5}_{-1.6}$	[85] SMY	04	FIT	SKAM + SNO
2.8–12.0	[86] AHMAD	02B	FIT	global solar
3.2–19.1	[87] FUKUDA	02	FIT	global solar

[74] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (ARAKI 05). *CPT* invariance is assumed.

[75] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using the data from Super-Kamiokande, SNO (AHMAD 02 and AHMAD 02B), and KamLAND (ARAKI 05) experiments. *CPT* invariance is assumed.

[76] HOSAKA 06 obtained this result by a two-neutrino oscillation analysis using the Super-Kamiokande and SNO (AHMAD 02 and AHMAD 02B) solar neutrino data.

[77] HOSAKA 06 obtained this result from the consistency between the observed and expected day-night flux asymmetry amplitude. The listed 68% CL range is derived from the 1σ boundary of the amplitude fit to the data. Oscillation parameters are constrained to be in the LMA region. The mixing angle is fixed at $\tan^2\theta = 0.44$ because the fit depends only very weekly on it.

[78] AHARMIM 05A obtained this result by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (ARAKI 05). *CPT* invariance is assumed. AHARMIM 05A also quotes $\Delta(m^2) = (8.0^{+0.6}_{-0.4}) \times 10^{-5}$ eV^2 as the error enveloping the 68% CL two-dimensional region.

[79] AHARMIM 05A obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in figure 35a of AHARMIM 05A. AHARMIM 05A also quotes $\Delta(m^2) = (6.5^{+4.4}_{-2.3}) \times 10^{-5}$ eV^2 as the error enveloping the 68% CL two-dimensional region.

[80] ARAKI 05 obtained this result by a two-neutrino oscillation analysis using KamLAND and solar neutrino data. *CPT* invariance is assumed. The 1σ error shown here is provided

by the KamLAND collaboration. The error quoted in ARAKI 05, $\Delta(m^2) = (7.9^{+0.6}_{-0.5}) \times 10^{-5}$, envelops the 68% CL two-dimensional region.

[81] AHMED 04A obtained this result by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (EGUCHI 03). *CPT* invariance is assumed. AHMED 04A also quotes $\Delta(m^2) = (7.1^{+1.2}_{-0.6}) \times 10^{-5}$ eV^2 as the error enveloping the 68% CL two-dimensional region.

[82] AHMED 04A obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 5(a) of AHMED 04A. The best-fit point is $\Delta(m^2) = 6.5 \times 10^{-5}$ eV^2, $\tan^2\theta = 0.40$ ($\sin^2 2\,\theta = 0.82$).

[83] SMY 04 obtained this result by a two-neutrino oscillation analysis using solar neutrino and KamLAND data (IANNI 03). *CPT* invariance is assumed.

[84] SMY 04 obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The 1σ errors are read from Fig. 6(a) of SMY 04.

[85] SMY 04 obtained this result by a two-neutrino oscillation analysis using the Super-Kamiokande and SNO (AHMAD 02 and AHMAD 02B) solar neutrino data. The 1σ errors are read from Fig. 6(a) of SMY 04.

[86] AHMAD 02B obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 4(b) of AHMAD 02B. The best fit point is $\Delta(m^2) = 5.0 \times 10^{-5}$ eV^2 and $\tan\theta = 0.34$ ($\sin^2 2\,\theta = 0.76$).

[87] FUKUDA 02 obtained this result by a two-neutrino oscillation analysis using the data from all solar neutrino experiments. The listed range of the parameter envelops the 95% CL two-dimensional region shown in Fig. 4 of FUKUDA 02. The best fit point is $\Delta(m^2) = 6.9 \times 10^{-5}$ eV^2 and $\tan^2\theta = 0.38$ ($\sin^2 2\,\theta = 0.80$).

$\sin^2(2\theta_{23})$

The ranges below correspond to the projection onto the $\sin^2(2\theta_{23})$ axis of the 90% CL contours in the $\sin^2(2\theta_{23}) - \Delta m^2_{32}$ plane presented by the authors.

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
>0.92	[88] ASHIE	05	SKAM	Super-Kamiokande
• • • We do not use the following data for averages, fits, limits, etc. • • •				
>0.2	[89] ADAMSON	06	MINS	atmospheric ν with far detector
>0.59	[90] AHN	06A	K2K	KEK to Super-K
>0.91	[91] HOSAKA	06A	SKAM	3ν oscillation; normal mass hierarchy
>0.86	[92] HOSAKA	06A	SKAM	3ν oscillation; inverted mass hierarchy
>0.7	[93] MICHAEL	06	MINS	MINOS
>0.58	[94] ALIU	05	K2K	KEK to Super-K
>0.6	[95] ALLISON	05	SOU2	
>0.80	[96] AMBROSIO	04	MCRO	MACRO
>0.90	[97] ASHIE	04	SKAM	L/E distribution
>0.30	[98] AHN	03	K2K	KEK to Super-K
>0.45	[99] AMBROSIO	03	MCRO	MACRO
>0.77	[100] AMBROSIO	03	MCRO	MACRO
>0.50	[101] SANCHEZ	03	SOU2	Soudan-2 Atmospheric
>0.80	[102] AMBROSIO	01	MCRO	upward μ
>0.82	[103] AMBROSIO	01	MCRO	upward μ

>0.45	[104] FUKUDA	99C	SKAM	upward μ
>0.70	[105] FUKUDA	99D	SKAM	upward μ
>0.30	[106] FUKUDA	99D	SKAM	stop μ / through
>0.82	[107] FUKUDA	98C	SKAM	Super-Kamiokande
>0.30	[108] HATAKEYAMA	98	KAMI	Kamiokande
>0.73	[109] HATAKEYAMA	98	KAMI	Kamiokande
>0.65	[110] FUKUDA	94	KAMI	Kamiokande

[88] ASHIE 05 obtained this result by a two-neutrino oscillation analysis using 92 kton yr atmospheric neutrino data from the complete Super-Kamiokande I running period.

[89] ADAMSON 06 obtained this result by a two-neutrino oscillation analysis of the L/E distribution using 4.54 kton yr atmospheric neutrino data with the MINOS far detector.

[90] Supercedes ALIU 05.

[91] HOSAKA 06A obtained this result ($\sin^2\theta_{23} = 0.37$–0.65) by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The normal mass hierarchy is assumed.

[92] HOSAKA 06A obtained this result ($\sin^2\theta_{23} = 0.37$–0.69) by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The inverted mass hierarchy is assumed.

[93] MICHAEL 06 best fit is for maximal mixing.

[94] The best fit is for maximal mixing.

[95] ALLISON 05 result is based upon atmospheric neutrino interactions including upward-stopping muons, with an exposure of 5.9 kton yr. From a two-flavor oscillation analysis the best-fit point is $\Delta m^2 = 0.0017$ eV^2 and $\sin^2(2\theta) = 0.97$.

[96] AMBROSIO 04 obtained this result, without using the absolute normalization of the neutrino flux, by combining the angular distribution of upward through-going muon tracks with $E_\mu > 1$ GeV, N_{low} and N_{high}, and the numbers of InDown + UpStop and InUp events. Here, N_{low} and N_{high} are the number of events with reconstructed neutrino energies $<$ 30 GeV and $>$ 130 GeV, respectively. InDown and InUp represent events with downward and upward-going tracks starting inside the detector due to neutrino interactions, while UpStop represents entering upward-going tracks which stop in the detector. The best fit is for maximal mixing.

[97] ASHIE 04 obtained this result from the L(flight length)/E(estimated neutrino energy) distribution of ν_μ disappearance probability, using the Super-Kamiokande-I 1489 live-day atmospheric neutrino data.

[98] There are several islands of allowed region from this K2K analysis, extending to high values of Δm^2. We only include the one that overlaps atmospheric neutrino analyses. The best fit is for maximal mixing.

[99] AMBROSIO 03 obtained this result on the basis of the ratio $R = N_{low}/N_{high}$, where N_{low} and N_{high} are the number of upward through-going muon events with reconstructed neutrino energy $<$ 30 GeV and $>$ 130 GeV, respectively. The data came from the full detector run started in 1994. The method of FELDMAN 98 is used to obtain the limits.

[100] AMBROSIO 03 obtained this result by using the ratio R and the angular distribution of the upward through-going muons. R is given in the previous note and the angular distribution is reported in AMBROSIO 01. The method of FELDMAN 98 is used to obtain the limits. The best fit is to maximal mixing.

[101] SANCHEZ 03 is based on an exposure of 5.9 kton yr. The result is obtained using a likelihood analysis of the neutrino L/E distribution for a selection μ flavor sample while the e-flavor sample provides flux normalization. The method of FELDMAN 98 is used to obtain the allowed region. The best fit is $\sin^2(2\theta) = 0.97$.

[102] AMBROSIO 01 result is based on the angular distribution of upward through-going muon tracks with $E_\mu > 1$ GeV. The data came from three different detector configurations, but the statistics is largely dominated by the full detector run, from May 1994 to December

2000. The total live time, normalized to the full detector configuration is 6.17 years. The best fit is obtained outside the physical region. The method of FELDMAN 98 is used to obtain the limits. The best fit is for maximal mixing.

[103] AMBROSIO 01 result is based on the angular distribution and normalization of upward through-going muon tracks with $E_\mu > 1$ GeV. See the previous footnote.

[104] FUKUDA 99C obtained this result from a total of 537 live days of upward through-going muon data in Super-Kamiokande between April 1996 to January 1998. With a threshold of $E_\mu > 1.6$ GeV, the observed flux is $(1.74 \pm 0.07 \pm 0.02) \times 10^{-13}$ $\text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is $\sin^2(2\theta) = 0.95$.

[105] FUKUDA 99D obtained this result from a simultaneous fitting to zenith angle distributions of upward-stopping and through-going muons. The flux of upward-stopping muons of minimum energy of 1.6 GeV measured between April 1996 and January 1998 is $(0.39 \pm 0.04 \pm 0.02) \times 10^{-13}$ $\text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. This is compared to the expected flux of $(0.73 \pm 0.16$ (theoretical error)$) \times 10^{-13}$ $\text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is to maximal mixing.

[106] FUKUDA 99D obtained this result from the zenith dependence of the upward-stopping/through-going flux ratio. The best fit is to maximal mixing.

[107] FUKUDA 98C obtained this result by an analysis of 33.0 kton yr atmospheric neutrino data. The best fit is for maximal mixing.

[108] HATAKEYAMA 98 obtained this result from a total of 2456 live days of upward-going muon data in Kamiokande between December 1985 and May 1995. With a threshold of $E_\mu > 1.6$ GeV, the observed flux of upward through-going muons is $(1.94 \pm 0.10 ^{+0.07}_{-0.06}) \times 10^{-13}$ $\text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. This is compared to the expected flux of $(2.46 \pm 0.54$ (theoretical error)$) \times 10^{-13}$ $\text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is for maximal mixing.

[109] HATAKEYAMA 98 obtained this result from a combined analysis of Kamiokande contained events (FUKUDA 94) and upward going muon events. The best fit is $\sin^2(2\theta) = 0.95$.

[110] FUKUDA 94 obtained the result by a combined analysis of sub- and multi-GeV atmospheric neutrino events in Kamiokande. The best fit is for maximal mixing.

Δm^2_{32}

The sign of Δm^2_{32} is not known at this time. Only the absolute value is quoted below. The ranges below correspond to the projection onto the Δm^2_{32} axis of the 90% CL contours in the $\sin^2(2\theta_{23}) - \Delta m^2_{32}$ plane presented by the authors.

VALUE (10^{-3} eV^2)	*DOCUMENT ID*		*TECN*	*COMMENT*
1.9 to 3.0	[111] ASHIE	04	SKAM	L/E distribution
• • • We do not use the following data for averages, fits, limits, etc. • • •				
0.07–50	[112] ADAMSON	06	MINS	atmospheric ν with far detector
1.9–4.0	[113,114] AHN	06A	K2K	KEK to Super-K
1.8–3.1	[115] HOSAKA	06A	SKAM	3ν oscillation; normal mass hierarchy
1.8–3.7	[116] HOSAKA	06A	SKAM	3ν oscillation; inverted mass hierarchy
2.2–3.8	[117] MICHAEL	06	MINS	MINOS
1.9–3.6	[113] ALIU	05	K2K	KEK to Super-K
0.3–12	[118] ALLISON	05	SOU2	
1.5–3.4	[119] ASHIE	05	SKAM	atmospheric neutrino
0.6–8.0	[120] AMBROSIO	04	MCRO	MACRO
1.5–3.9	[121] AHN	03	K2K	KEK to Super-K
0.25–9.0	[122] AMBROSIO	03	MCRO	MACRO

0.6–7.0	[123] AMBROSIO	03	MCRO	MACRO
0.15–15	[124] SANCHEZ	03	SOU2	Soudan-2 Atmospheric
0.6–15	[125] AMBROSIO	01	MCRO	upward μ
1.0–6.0	[126] AMBROSIO	01	MCRO	upward μ
1.0–50	[127] FUKUDA	99C	SKAM	upward μ
1.5–15.0	[128] FUKUDA	99D	SKAM	upward μ
0.7–18	[129] FUKUDA	99D	SKAM	stop μ / through
0.5–6.0	[130] FUKUDA	98C	SKAM	Super-Kamiokande
0.55–50	[131] HATAKEYAMA	98	KAMI	Kamiokande
4–23	[132] HATAKEYAMA	98	KAMI	Kamiokande
5–25	[133] FUKUDA	94	KAMI	Kamiokande

[111] ASHIE 04 obtained this result from the L(flight length)/E(estimated neutrino energy) distribution of ν_μ disappearance probability, using the Super-Kamiokande-I 1489 live-day atmospheric neutrino data. The best fit is for $\Delta m^2 = 2.4 \times 10^{-3}$ eV^2.

[112] ADAMSON 06 obtained this result by a two-neutrino oscillation analysis of the L/E distribution using 4.54 kton yr atmospheric neutrino data with the MINOS far detector.

[113] The best fit in the physical region is for $\Delta m^2 = 2.8 \times 10^{-3}$ eV^2.

[114] Supercedes ALIU 05.

[115] HOSAKA 06A obtained this result by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The normal mass hierarchy is assumed.

[116] HOSAKA 06A obtained this result by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The inverted mass hierarchy is assumed.

[117] MICHAEL 06 best fit is 2.74×10^{-3} eV^2.

[118] ALLISON 05 result is based on an atmospheric neutrino observation with an exposure of 5.9 kton yr. From a two-flavor oscillation analysis the best-fit point is $\Delta m^2 = 0.0017$ eV^2 and $\sin^2 2\,\theta = 0.97$.

[119] ASHIE 05 obtained this result by a two-neutrino oscillation analysis using 92 kton yr atmospheric neutrino data from the complete Super-Kamiokande I running period. The best fit is for $\Delta m^2 = 2.1 \times 10^{-3}$ eV^2.

[120] AMBROSIO 04 obtained this result, without using the absolute normalization of the neutrino flux, by combining the angular distribution of upward through-going muon tracks with $E_\mu > 1$ GeV, N_{low} and N_{high}, and the numbers of InDown + UpStop and InUp events. Here, N_{low} and N_{high} are the number of events with reconstructed neutrino energies < 30 GeV and > 130 GeV, respectively. InDown and InUp represent events with downward and upward-going tracks starting inside the detector due to neutrino interactions, while UpStop represents entering upward-going tracks which stop in the detector. The best fit is for $\Delta m^2 = 2.3 \times 10^{-3}$ eV^2.

[121] There are several islands of allowed region from this K2K analysis, extending to high values of Δm^2. We only include the one that overlaps atmospheric neutrino analyses. The best fit is for $\Delta m^2 = 2.8 \times 10^{-3}$ eV^2.

[122] AMBROSIO 03 obtained this result on the basis of the ratio $R = N_{low}/N_{high}$, where N_{low} and N_{high} are the number of upward through-going muon events with reconstructed neutrino energy < 30 GeV and > 130 GeV, respectively. The data came from the full detector run started in 1994. The method of FELDMAN 98 is used to obtain the limits. The best fit is for $\Delta m^2 = 2.5 \times 10^{-3}$ eV^2.

[123] AMBROSIO 03 obtained this result by using the ratio R and the angular distribution of the upward through-going muons. R is given in the previous note and the angular distribution is reported in AMBROSIO 01. The method of FELDMAN 98 is used to obtain the limits. The best fit is for $\Delta m^2 = 2.5 \times 10^{-3}$ eV^2.

[124] SANCHEZ 03 is based on an exposure of 5.9 kton yr. The result is obtained using a likelihood analysis of the neutrino L/E distribution for a selection μ flavor sample while

the *e*-flavor sample provides flux normalization. The method of FELDMAN 98 is used to obtain the allowed region. The best fit is for $\Delta m^2 = 5.2 \times 10^{-3}\ \text{eV}^2$.

[125] AMBROSIO 01 result is based on the angular distribution of upward through-going muon tracks with $E_\mu > 1$ GeV. The data came from three different detector configurations, but the statistics is largely dominated by the full detector run, from May 1994 to December 2000. The total live time, normalized to the full detector configuration is 6.17 years. The best fit is obtained outside the physical region. The method of FELDMAN 98 is used to obtain the limits.

[126] AMBROSIO 01 result is based on the angular distribution and normalization of upward through-going muon tracks with $E_\mu > 1$ GeV. See the previous footnote.

[127] FUKUDA 99C obtained this result from a total of 537 live days of upward through-going muon data in Super-Kamiokande between April 1996 to January 1998. With a threshold of $E_\mu > 1.6$ GeV, the observed flux is $(1.74 \pm 0.07 \pm 0.02) \times 10^{-13}\ \text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is for $\Delta m^2 = 5.9 \times 10^{-3}\ \text{eV}^2$.

[128] FUKUDA 99D obtained this result from a simultaneous fitting to zenith angle distributions of upward-stopping and through-going muons. The flux of upward-stopping muons of minimum energy of 1.6 GeV measured between April 1996 and January 1998 is $(0.39 \pm 0.04 \pm 0.02) \times 10^{-13}\ \text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. This is compared to the expected flux of $(0.73 \pm 0.16$ (theoretical error)$) \times 10^{-13}\ \text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is for $\Delta m^2 = 3.9 \times 10^{-3}\ \text{eV}^2$.

[129] FUKUDA 99D obtained this result from the zenith dependence of the upward-stopping/through-going flux ratio. The best fit is for $\Delta m^2 = 3.1 \times 10^{-3}\ \text{eV}^2$.

[130] FUKUDA 98C obtained this result by an analysis of 33.0 kton yr atmospheric neutrino data. The best fit is for $\Delta m^2 = 2.2 \times 10^{-3}\ \text{eV}^2$.

[131] HATAKEYAMA 98 obtained this result from a total of 2456 live days of upward-going muon data in Kamiokande between December 1985 and May 1995. With a threshold of $E_\mu > 1.6$ GeV, the observed flux of upward through-going muons is $(1.94 \pm 0.10^{+0.07}_{-0.06}) \times 10^{-13}\ \text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. This is compared to the expected flux of $(2.46 \pm 0.54$ (theoretical error)$) \times 10^{-13}\ \text{cm}^{-2}\text{s}^{-1}\text{sr}^{-1}$. The best fit is for $\Delta m^2 = 2.2 \times 10^{-3}\ \text{eV}^2$.

[132] HATAKEYAMA 98 obtained this result from a combined analysis of Kamiokande contained events (FUKUDA 94) and upward going muon events. The best fit is for $\Delta m^2 = 13 \times 10^{-3}\ \text{eV}^2$.

[133] FUKUDA 94 obtained the result by a combined analysis of sub- and multi-GeV atmospheric neutrino events in Kamiokande. The best fit is for $\Delta m^2 = 16 \times 10^{-3}\ \text{eV}^2$.

$\sin^2(2\theta_{13})$

At present time, limits of $\sin^2(2\ \theta_{13})$ are derived from the search for the reactor $\overline{\nu}_e$ disappearance at distances corresponding to the Δm^2_{23} value, i.e. L $\sim$ 1km. Alternatively, somewhat weaker limits can be obtained from the analysis of the solar neutrino data.

VALUE	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
<0.19	90	[134] APOLLONIO 99	CHOZ	Reactor Experiment

• • • We do not use the following data for averages, fits, limits, etc. • • •

VALUE	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
<0.48	90	[135] HOSAKA 06A	SKAM	3ν oscillation; normal mass hierarchy
<0.79	90	[136] HOSAKA 06A	SKAM	3ν oscillation; inverted mass hierarchy
<0.36		[137] YAMAMOTO 06	K2K	Accelerator experiment
<0.48	90	[138] AHN 04	K2K	Accelerator experiment
<0.36	90	[139] BOEHM 01		Palo Verde react.
<0.45	90	[140] BOEHM 00		Palo Verde react.

[134] The quoted limit is for $\Delta m^2_{32} = 1.9 \times 10^{-3}$ eV^2. That value of Δm^2_{32} is the 1-σ low value for ALIU 05. For the ALIU 05 best fit value of 2.8×10^{-3} eV^2, the $\sin^2 2\,\theta_{13}$ limit is < 0.13. See also APOLLONIO 03 for a detailed description of the experiment.

[135] HOSAKA 06A obtained this result by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The normal mass hierarchy is assumed.

[136] HOSAKA 06A obtained this result by a three-neutrino oscillation analysis with one mass scale dominance ($\Delta m^2_{21} = 0$) using the Super-Kamiokande-I atmospheric neutrino data. The inverted mass hierarchy is assumed.

[137] YAMAMOTO 06 searched for $\nu_\mu \rightarrow \nu_e$ appearance. Assumes $2\sin^2(2\theta_{\mu e}) = \sin^2(2\theta_{13})$. The quoted limit is for $\Delta m^2_{32} = 1.9 \times 10^{-3}$ eV^2. That value of Δm^2_{32} is the one-σ low value for AHN 06A. For the AHN 06A best fit value of 2.8×10^{-3} eV^2, the $\sin^2(2\theta_{13})$ limit is < 0.26. Supersedes AHN 04.

[138] AHN 04 searched for $\nu_\mu \rightarrow \nu_e$ appearance. Assuming $2\sin^2(2\,\theta_{\mu_e}) = \sin^2(2\,\theta_{13})$, a limit on $\sin^2(2\,\theta_{\mu_e})$ is converted to a limit on $\sin^2(2\,\theta_{13})$.The quoted limit is for $\Delta m^2_{32} = 1.9 \times 10^{-3}$ eV^2. That value of Δm^2_{32} is the one-σ low value for ALIU 05. For the ALIU 05 best fit value of 2.8×10^{-3} eV^2, the $\sin^2(2\,\theta_{13})$ limit is < 0.30.

[139] The quoted limit is for $\Delta m^2_{32} = 1.9 \times 10^{-3}$ eV^2. That value of Δm^2_{32} is the 1-σ low value for ALIU 05. For the ALIU 05 best fit value of 2.8×10^{-3} eV^2, the $\sin^2 2\,\theta_{13}$ limit is < 0.19. In this range, the θ_{13} limit is larger for lower values of Δm^2_{32}, and smaller for higher values of Δm^2_{32}.

[140] The quoted limit is for $\Delta m^2_{32} = 1.9 \times 10^{-3}$ eV^2. That value of Δm^2_{32} is the 1-σ low value for ALIU 05. For the ALIU 05 best fit value of 2.8×10^{-3} eV^2, the $\sin^2 2\,\theta_{13}$ limit is < 0.23.

(C) Other neutrino mixing results

The LSND collaboration reported in AGUILAR 01 a signal which is consistent with $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ oscillations. In a three neutrino framework, this would be a measurement of θ_{12} and Δm^2_{21}. This does not appear to be consistent with the interpretation of other neutrino data, particularly solar neutrino experiments. If the LSND anomaly is correct, a more complicated framework is required, perhaps involving one or more sterile neutrinos, or even CPT violation. The following listings include results which might be relevant towards understanding or ruling out the LSND observations. They include searches for $\nu_\mu \rightarrow \nu_e$, $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$, sterile neutrino oscillations, and *CPT* violation.

$\Delta(m^2)$ for $\sin^2(2\theta) = 1$ $(\nu_\mu \rightarrow \nu_e)$

VALUE (eV^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •					
<0.0008	90	AHN	04	K2K	Water Cherenkov
<0.4	90	ASTIER	03	NOMD	CERN SPS
<2.4	90	AVVAKUMOV	02	NTEV	NUTEV FNAL
		[141] AGUILAR	01	LSND	$\nu\mu \rightarrow \nu_e$ osc.prob.

0.03 to 0.3	95	[142] ATHANASSO...98	LSND	$\nu_\mu \rightarrow \nu_e$
<2.3	90	[143] LOVERRE 96		CHARM/CDHS
<0.9	90	VILAIN 94C	CHM2	CERN SPS
<0.09	90	ANGELINI 86	HLBC	BEBC CERN PS

[141] AGUILAR 01 is the final analysis of the LSND full data set. Search is made for the $\nu_\mu \rightarrow \nu_e$ oscillations using ν_μ from π^+ decay in flight by observing beam-on electron events from $\nu_e \mathrm{C} \rightarrow e^- X$. Present analysis results in 8.1 ± 12.2 ± 1.7 excess events in the $60 < E_e < 200$ MeV energy range, corresponding to oscillation probability of $0.10 \pm 0.16 \pm 0.04\%$. This is consistent, though less significant, with the previous result of ATHANASSOPOULOS 98, which it supersedes. The present analysis uses selection criteria developed for the decay at rest region, and is less effective in removing the background above 60 MeV than ATHANASSOPOULOS 98.

[142] ATHANASSOPOULOS 98 is a search for the $\nu_\mu \rightarrow \nu_e$ oscillations using ν_μ from π^+ decay in flight. The 40 observed beam-on electron events are consistent with $\nu_e \mathrm{C} \rightarrow e^- \mathrm{X}$; the expected background is 21.9 ± 2.1. Authors interpret this excess as evidence for an oscillation signal corresponding to oscillations with probability $(0.26 \pm 0.10 \pm 0.05)\%$. Although the significance is only $2.3\,\sigma$, this measurement is an important and consistent cross check of ATHANASSOPOULOS 96 who reported evidence for $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ oscillations from μ^+ decay at rest. See also ATHANASSOPOULOS 98B.

[143] LOVERRE 96 uses the charged-current to neutral-current ratio from the combined CHARM (ALLABY 86) and CDHS (ABRAMOWICZ 86) data from 1986.

$\sin^2(2\theta)$ for "Large" $\Delta(m^2)$ $(\nu_\mu \rightarrow \nu_e)$

VALUE (units 10^{-3})	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<110	90	[144] AHN 04	K2K	Water Cherenkov
< 1.4	90	ASTIER 03	NOMD	CERN SPS
< 1.6	90	AVVAKUMOV 02	NTEV	NUTEV FNAL
		[145] AGUILAR 01	LSND	$\nu\mu \rightarrow \nu_e$ osc.prob.
0.5 to 30	95	[146] ATHANASSO...98	LSND	$\nu_\mu \rightarrow \nu_e$
< 3.0	90	[147] LOVERRE 96		CHARM/CDHS
< 9.4	90	VILAIN 94C	CHM2	CERN SPS
< 5.6	90	[148] VILAIN 94C	CHM2	CERN SPS

[144] The limit becomes $\sin^2 2\theta < 0.15$ at $\Delta m^2 = 2.8 \times 10^{-3}$ eV^2, the bets-fit value of the ν_μ disappearance analysis in K2K.

[145] AGUILAR 01 is the final analysis of the LSND full data set of the search for the $\nu_\mu \rightarrow \nu_e$ oscillations. See footnote in preceding table for further details.

[146] ATHANASSOPOULOS 98 report $(0.26 \pm 0.10 \pm 0.05)\%$ for the oscillation probability; the value of $\sin^2 2\theta$ for large Δm^2 is deduced from this probability. See footnote in preceding table for further details, and see the paper for a plot showing allowed regions. If effect is due to oscillation, it is most likely to be intermediate $\sin^2 2\theta$ and Δm^2. See also ATHANASSOPOULOS 98B.

[147] LOVERRE 96 uses the charged-current to neutral-current ratio from the combined CHARM (ALLABY 86) and CDHS (ABRAMOWICZ 86) data from 1986.

[148] VILAIN 94C limit derived by combining the ν_μ and $\overline{\nu}_\mu$ data assuming *CP* conservation.

$\Delta(m^2)$ for $\sin^2(2\theta) = 1$ $(\overline{\nu}_\mu \rightarrow \overline{\nu}_e)$

VALUE (eV^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •					
<0.055	90	[149] ARMBRUSTER	02	KAR2	Liquid Sci. calor.
<2.6	90	AVVAKUMOV	02	NTEV	NUTEV FNAL
0.03–0.05		[150] AGUILAR	01	LSND	LAMPF
0.05–0.08	90	[151] ATHANASSO...	96	LSND	LAMPF
0.048–0.090	80	[152] ATHANASSO...	95		
<0.07	90	[153] HILL	95		
<0.9	90	VILAIN	94C	CHM2	CERN SPS
<0.14	90	[154] FREEDMAN	93	CNTR	LAMPF

[149] ARMBRUSTER 02 is the final analysis of the KARMEN 2 data for 17.7 m distance from the ISIS stopped pion and muon neutrino source. It is a search for $\overline{\nu}_e$, detected by the inverse β-decay reaction on protons and ^{12}C. 15 candidate events are observed, and 15.8 ± 0.5 background events are expected, hence no oscillation signal is detected. The results exclude large regions of the parameter area favored by the LSND experiment.

[150] AGUILAR 01 is the final analysis of the LSND full data set. It is a search for $\overline{\nu}_e$ 30 m from LAMPF beam stop. Neutrinos originate mainly for π^+ decay at rest. $\overline{\nu}_e$ are detected through $\overline{\nu}_e p \rightarrow e^+ n$ ($20 < E_{e^+} < 60$ MeV) in delayed coincidence with $np \rightarrow d\gamma$. AUthors observe 87.9 ± 22.4 ± 6.0 total excess events. The observation is attributed to $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ oscillations with the oscillation probability of 0.264 ± 0.067 ± 0.045%, consistent with the previously published result. Taking into account all constraints, the most favored allowed region of oscillation parameters is a band of $\Delta(m^2)$ from 0.2–2.0 eV^2. Supersedes ATHANASSOPOULOS 95, ATHANASSOPOULOS 96, and ATHANASSOPOULOS 98.

[151] ATHANASSOPOULOS 96 is a search for $\overline{\nu}_e$ 30 m from LAMPF beam stop. Neutrinos originate mainly from π^+ decay at rest. $\overline{\nu}_e$ could come from either $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ or $\nu_e \rightarrow \overline{\nu}_e$; our entry assumes the first interpretation. They are detected through $\overline{\nu}_e p \rightarrow e^+ n$ (20 MeV $< E_{e^+} <$ 60 MeV) in delayed coincidence with $np \rightarrow d\gamma$. Authors observe 51 ± 20 ± 8 total excess events over an estimated background 12.5 ± 2.9. ATHANASSOPOULOS 96B is a shorter version of this paper.

[152] ATHANASSOPOULOS 95 error corresponds to the 1.6σ band in the plot. The expected background is 2.7 ± 0.4 events. Corresponds to an oscillation probability of $(0.34^{+0.20}_{-0.18} \pm 0.07)$%. For a different interpretation, see HILL 95. Replaced by ATHANASSOPOULOS 96.

[153] HILL 95 is a report by one member of the LSND Collaboration, reporting a different conclusion from the analysis of the data of this experiment (see ATHANASSOPOULOS 95). Contrary to the rest of the LSND Collaboration, Hill finds no evidence for the neutrino oscillation $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ and obtains only upper limits.

[154] FREEDMAN 93 is a search at LAMPF for $\overline{\nu}_e$ generated from any of the three neutrino types ν_μ, $\overline{\nu}_\mu$, and ν_e which come from the beam stop. The $\overline{\nu}_e$'s would be detected by the reaction $\overline{\nu}_e p \rightarrow e^+ n$. FREEDMAN 93 replaces DURKIN 88.

$\sin^2(2\theta)$ for "Large" $\Delta(m^2)$ $(\overline{\nu}_\mu \rightarrow \overline{\nu}_e)$

VALUE (units 10^{-3})	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •					
<1.7	90	[155] ARMBRUSTER	02	KAR2	Liquid Sci. calor.
<1.1	90	AVVAKUMOV	02	NTEV	NUTEV FNAL
5.3±1.3±9.0		[156] AGUILAR	01	LSND	LAMPF

$6.2 \pm 2.4 \pm 1.0$		[157] ATHANASSO...96	LSND	LAMPF
3–12	80	[158] ATHANASSO...95		
<6	90	[159] HILL 95		

[155] ARMBRUSTER 02 is the final analysis of the KARMEN 2 data. See footnote in the preceding table for further details, and the paper for the exclusion plot.

[156] AGUILAR 01 is the final analysis of the LSND full data set. The deduced oscillation probability is $0.264 \pm 0.067 \pm 0.045\%$; the value of $\sin^2 2\theta$ for large $\Delta(m^2)$ is twice this probability (although these values are excluded by other constraints). See footnote in preceding table for further details, and the paper for a plot showing allowed regions. Supersedes ATHANASSOPOULOS 95, ATHANASSOPOULOS 96, and ATHANASSOPOULOS 98.

[157] ATHANASSOPOULOS 96 reports $(0.31 \pm 0.12 \pm 0.05)\%$ for the oscillation probability; the value of $\sin^2 2\theta$ for large $\Delta(m^2)$ should be twice this probability. See footnote in preceding table for further details, and see the paper for a plot showing allowed regions.

[158] ATHANASSOPOULOS 95 error corresponds to the 1.6σ band in the plot. The expected background is 2.7 ± 0.4 events. Corresponds to an oscillation probability of $(0.34^{+0.20}_{-0.18} \pm 0.07)\%$. For a different interpretation, see HILL 95. Replaced by ATHANASSOPOULOS 96.

[159] HILL 95 is a report by one member of the LSND Collaboration, reporting a different conclusion from the analysis of the data of this experiment (see ATHANASSOPOULOS 95). Contrary to the rest of the LSND Collaboration, Hill finds no evidence for the neutrino oscillation $\overline{\nu}_\mu \rightarrow \overline{\nu}_e$ and obtains only upper limits.

$\Delta(m^2)$ for $\sin^2(2\theta) = 1$ $\quad (\nu_\mu(\overline{\nu}_\mu) \rightarrow \nu_e(\overline{\nu}_e))$

VALUE (eV^2)	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
<0.075	90	BORODOV... 92	CNTR	BNL E776
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<1.6	90	[160] ROMOSAN 97	CCFR	FNAL

[160] ROMOSAN 97 uses wideband beam with a 0.5 km decay region.

$\sin^2(2\theta)$ for "Large" $\Delta(m^2)$ $\quad (\nu_\mu(\overline{\nu}_\mu) \rightarrow \nu_e(\overline{\nu}_e))$

VALUE (units 10^{-3})	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
<1.8	90	[161] ROMOSAN 97	CCFR	FNAL
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<3.8	90	[162] MCFARLAND 95	CCFR	FNAL
<3	90	BORODOV... 92	CNTR	BNL E776

[161] ROMOSAN 97 uses wideband beam with a 0.5 km decay region.

[162] MCFARLAND 95 state that "This result is the most stringent to date for $250 < \Delta(m^2) < 450$ eV^2 and also excludes at 90%CL much of the high $\Delta(m^2)$ region favored by the recent LSND observation." See ATHANASSOPOULOS 95 and ATHANASSOPOULOS 96.

Sterile neutrino limits from atmospheric neutrino studies

$\Delta(m^2)$ for $\sin^2(2\theta) = 1$ $(\nu_\mu \rightarrow \nu_s)$

ν_s means ν_τ or any sterile (noninteracting) ν.

VALUE (10^{-5} eV^2)	*CL%*	*DOCUMENT ID*	*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
<3000 (or <550)	90	[163] OYAMA 89	KAMI	Water Cherenkov
< 4.2 or > 54.	90	BIONTA 88	IMB	Flux has ν_μ, $\overline{\nu}_\mu$, ν_e, and $\overline{\nu}_e$

[163] OYAMA 89 gives a range of limits, depending on assumptions in their analysis. They argue that the region $\Delta(m^2) = (100\text{–}1000) \times 10^{-5}$ eV^2 is not ruled out by any data for large mixing.

Search for $\nu_\mu \to \nu_s$

VALUE	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •				
	[164] AMBROSIO	01	MCRO	matter effects
	[165] FUKUDA	00	SKAM	neutral currents + matter effects

[164] AMBROSIO 01 tested the pure 2-flavor $\nu_\mu \to \nu_s$ hypothesis using matter effects which change the shape of the zenith-angle distribution of upward through-going muons. With maximum mixing and $\Delta(m^2)$ around 0.0024 eV^2, the $\nu_\mu \to \nu_s$ oscillation is disfavored with 99% confidence level with respect to the $\nu_\mu \to \nu_\tau$ hypothesis.

[165] FUKUDA 00 tested the pure 2-flavor $\nu_\mu \to \nu_s$ hypothesis using three complementary atmospheric-neutrino data samples. With this hypothesis, zenith-angle distributions are expected to show characteristic behavior due to neutral currents and matter effects. In the $\Delta(m^2)$ and $\sin^2 2\theta$ region preferred by the Super-Kamiokande data, the $\nu_\mu \to \nu_s$ hypothesis is rejected at the 99% confidence level, while the $\nu_\mu \to \nu_\tau$ hypothesis consistently fits all of the data sample.

——— *CPT* tests ———

$\langle \Delta m^2_{21} - \Delta \overline{m}^2_{21} \rangle$

VALUE (10^{-4} eV^2)	*CL%*	*DOCUMENT ID*		*TECN*	*COMMENT*
• • • We do not use the following data for averages, fits, limits, etc. • • •					
<1.1	99.7	[166] DEGOUVEA	05	FIT	solar vs. reactor

[166] DEGOUVEA 05 obtained this bound at the 3σ CL from the KamLAND (ARAKI 05) and solar neutrino data.

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SMY	04	PR D69 011104R	M.B. Smy *et al.*	(Super-Kamiokande Collab.)
AHN	03	PRL 90 041801	M.H. Ahn *et al.*	(K2K Collab.)
AMBROSIO	03	PL B566 35	M. Ambrosio *et al.*	(MACRO Collab.)
APOLLONIO	03	EPJ C27 331	M. Apollonio *et al.*	(CHOOZ Collab.)
ASTIER	03	PL B570 19	P. Astier *et al.*	(NOMAD Collab.)

EGUCHI	03	PRL 90 021802	K. Eguchi *et al.*	(KamLAND Collab.)
GANDO	03	PRL 90 171302	Y. Gando *et al.*	(Super-Kamiokande Collab.)
IANNI	03	JPG 29 2107	A. Ianni	(INFN Gran Sasso)
SANCHEZ	03	PR D68 113004	M. Sanchez *et al.*	(Soudan 2 Collab.)
ABDURASHI...	02	JETP 95 181	J.N. Abdurashitov *et al.*	(SAGE Collab.)
		Translated from ZETF 122 211.		
AHMAD	02	PRL 89 011301	Q.R. Ahmad *et al.*	(SNO Collab.)
AHMAD	02B	PRL 89 011302	Q.R. Ahmad *et al.*	(SNO Collab.)
ARMBRUSTER	02	PR D65 112001	B. Armbruster *et al.*	(KARMEN 2 Collab.)
AVVAKUMOV	02	PRL 89 011804	S. Avvakumov *et al.*	(NuTeV Collab.)
FUKUDA	02	PL B539 179	S. Fukuda *et al.*	(Super-Kamiokande Collab.)
AGUILAR	01	PR D64 112007	A. Aguilar *et al.*	(LSND Collab.)
AHMAD	01	PRL 87 071301	Q.R. Ahmad *et al.*	(SNO Collab.)
AMBROSIO	01	PL B517 59	M. Ambrosio *et al.*	(MACRO Collab.)
BOEHM	01	PR D64 112001	F. Boehm *et al.*	
FUKUDA	01	PRL 86 5651	S. Fukuda *et al.*	(Super-Kamiokande Collab.)
AMBROSIO	00	PL B478 5	M. Ambrosio *et al.*	(MACRO Collab.)
BOEHM	00	PRL 84 3764	F. Boehm *et al.*	
FUKUDA	00	PRL 85 3999	S. Fukuda *et al.*	(Super-Kamiokande Collab.)
ABDURASHI...	99B	PR C60 055801	J.N. Abdurashitov *et al.*	(SAGE Collab.)
ALLISON	99	PL B449 137	W.W.M. Allison *et al.*	(Soudan 2 Collab.)
APOLLONIO	99	PL B466 415	M. Apollonio *et al.*	(CHOOZ Collab.)
Also		PL B472 434 (erratum)	M. Apollonio *et al.*	(CHOOZ Collab.)
FUKUDA	99C	PRL 82 2644	Y. Fukuda *et al.*	(Super-Kamiokande Collab.)
FUKUDA	99D	PL B467 185	Y. Fukuda *et al.*	(Super-Kamiokande Collab.)
HAMPEL	99	PL B447 127	W. Hampel *et al.*	(GALLEX Collab.)
AMBROSIO	98	PL B434 451	M. Ambrosio *et al.*	(MACRO Collab.)
APOLLONIO	98	PL B420 397	M. Apollonio *et al.*	(CHOOZ Collab.)
ATHANASSO...	98	PRL 81 1774	C. Athanassopoulos *et al.*	(LSND Collab.)
ATHANASSO...	98B	PR C58 2489	C. Athanassopoulos *et al.*	(LSND Collab.)
CLEVELAND	98	APJ 496 505	B.T. Cleveland *et al.*	(Homestake Collab.)
FELDMAN	98	PR D57 3873	G.J. Feldman, R.D. Cousins	
FUKUDA	98C	PRL 81 1562	Y. Fukuda *et al.*	(Super-Kamiokande Collab.)
HATAKEYAMA	98	PRL 81 2016	S. Hatakeyama *et al.*	(Kamiokande Collab.)
CLARK	97	PRL 79 345	R. Clark *et al.*	(IMB Collab.)
ROMOSAN	97	PRL 78 2912	A. Romosan *et al.*	(CCFR Collab.)
AGLIETTA	96	JETPL 63 791	M. Aglietta *et al.*	(LSD Collab.)
		Translated from ZETFP 63 753.		
ATHANASSO...	96	PR C54 2685	C. Athanassopoulos *et al.*	(LSND Collab.)
ATHANASSO...	96B	PRL 77 3082	C. Athanassopoulos *et al.*	(LSND Collab.)
FUKUDA	96	PRL 77 1683	Y. Fukuda *et al.*	(Kamiokande Collab.)
FUKUDA	96B	PL B388 397	Y. Fukuda *et al.*	(Kamiokande Collab.)
GREENWOOD	96	PR D53 6054	Z.D. Greenwood *et al.*	(UCI, SVR, SCUC)
HAMPEL	96	PL B388 384	W. Hampel *et al.*	(GALLEX Collab.)
LOVERRE	96	PL B370 156	P.F. Loverre	
ACHKAR	95	NP B434 503	B. Achkar *et al.*	(SING, SACLD, CPPM, CDEF+)
AHLEN	95	PL B357 481	S.P. Ahlen *et al.*	(MACRO Collab.)
ATHANASSO...	95	PRL 75 2650	C. Athanassopoulos *et al.*	(LSND Collab.)
DAUM	95	ZPHY C66 417	K. Daum *et al.*	(FREJUS Collab.)
HILL	95	PRL 75 2654	J.E. Hill	(PENN)
MCFARLAND	95	PRL 75 3993	K.S. McFarland *et al.*	(CCFR Collab.)
DECLAIS	94	PL B338 383	Y. Declais *et al.*	
FUKUDA	94	PL B335 237	Y. Fukuda *et al.*	(Kamiokande Collab.)
VILAIN	94C	ZPHY C64 539	P. Vilain *et al.*	(CHARM II Collab.)
FREEDMAN	93	PR D47 811	S.J. Freedman *et al.*	(LAMPF E645 Collab.)
BECKER-SZ...	92B	PR D46 3720	R.A. Becker-Szendy *et al.*	(IMB Collab.)
BEIER	92	PL B283 446	E.W. Beier *et al.*	(KAM2 Collab.)
Also		PTRSL A346 63	E.W. Beier, E.D. Frank	(PENN)
BORODOV...	92	PRL 68 274	L. Borodovsky *et al.*	(COLU, JHU, ILL)
HIRATA	92	PL B280 146	K.S. Hirata *et al.*	(Kamiokande II Collab.)
CASPER	91	PRL 66 2561	D. Casper *et al.*	(IMB Collab.)
HIRATA	91	PRL 66 9	K.S. Hirata *et al.*	(Kamiokande II Collab.)
KUVSHINN...	91	JETPL 54 253	A.A. Kuvshinnikov *et al.*	(KIAE)
BERGER	90B	PL B245 305	C. Berger *et al.*	(FREJUS Collab.)
HIRATA	90	PRL 65 1297	K.S. Hirata *et al.*	(Kamiokande II Collab.)
AGLIETTA	89	EPL 8 611	M. Aglietta *et al.*	(FREJUS Collab.)
DAVIS	89	ARNPS 39 467	R. Davis, A.K. Mann, L. Wolfenstein	(BNL, PENN+)
OYAMA	89	PR D39 1481	Y. Oyama *et al.*	(Kamiokande II Collab.)
BIONTA	88	PR D38 768	R.M. Bionta *et al.*	(IMB Collab.)
DURKIN	88	PRL 61 1811	L.S. Durkin *et al.*	(OSU, ANL, CIT+)
ABRAMOWICZ	86	PRL 57 298	H. Abramowicz *et al.*	(CDHS Collab.)

ALLABY	86	PL B177 446	J.V. Allaby *et al.*	(CHARM Collab.)
ANGELINI	86	PL B179 307	C. Angelini *et al.*	(PISA, ATHU, PADO+)
VUILLEUMIER	82	PL 114B 298	J.L. Vuilleumier *et al.*	(CIT, SIN, MUNI)
BOLIEV	81	SJNP 34 787	M.M. Boliev *et al.*	(INRM)
		Translated from YAF 34 1418.		
KWON	81	PR D24 1097	H. Kwon *et al.*	(CIT, ISNG, MUNI)
BOEHM	80	PL 97B 310	F. Boehm *et al.*	(ILLG, CIT, ISNG, MUNI)
CROUCH	78	PR D18 2239	M.F. Crouch *et al.*	(CASE, UCI, WITW)